

NEUTRINOS (Lecture 1 of 3)

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Tentative Outline for The Lectures

1. Brief History of the Neutrino;
2. Neutrino Puzzles – The Discovery of Neutrino Masses;
3. Neutrino Oscillations;
4. What We Know We Don't Know;
5. Neutrino Masses As Physics Beyond the Standard Model;
6. Some Ideas for Tiny Neutrino Masses, and Some Consequences.

[note: Questions/Suggestions/Complaints are ALWAYS welcome]

Some Neutrino references (WARNING: Biased Sample)

- “Are There Really Neutrinos? – An Evidential History,” Allan Franklin, Perseus Books, 2001. Good discussion of neutrino history.
- A. de Gouvêa, “TASI lectures on neutrino physics,” hep-ph/0411274;
- R. N. Mohapatra, A. Yu. Smirnov, “Neutrino Mass and New Physics,” Ann. Rev. Nucl. Part. Sci. **56**, 569 (2006) [hep-ph/0603118];
- M. C. Gonzalez-Garcia, M. Maltoni, “Phenomenology with Massive Neutrinos,” Phys. Rept. **460**, 1 (2008) [arXiv:0704.1800 [hep-ph]];
- C. Giunti and C.W. Kim, “Fundamentals of Neutrino Physics and Astrophysics,” Oxford University Press (2007);
- “The Physics of Neutrinos,” V. Barger, D. Marfatia, K. Whisnant, Princeton University Press (2012);
- A. de Gouvêa *et al.*, “Working Group Report: Neutrinos,” arXiv:1310.4340;
- A. de Gouvêa, “Neutrino Mass Models,” Ann. Rev. Nucl. Part. Sci. **66**, 197 (2016).
- Several lectures at TASI 2020; soon in an arXiv near you.

1 - Brief History of the Neutrino

1. 1896: Henri Becquerel discovers natural radioactivity while studying phosphorescent properties of uranium salts.
 - α rays: easy to absorb, hard to bend, positive charge, mono-energetic;
 - β rays: harder to absorb, easy to bend, negative charge, spectrum?;
 - γ rays: no charge, very hard to absorb.
2. 1897: J.J. Thompson discovers the electron.
3. 1914: Chadwick presents definitive evidence for a continuous β -ray spectrum. Origin unknown. Different options include several different energy loss mechanisms.

It took 15+ years to decide that the “real” β -ray spectrum was really continuous. Reason for continuous spectrum was a total mystery:

- QM: Spectra are discrete;
- Energy-momentum conservation: $N \rightarrow N' + e^-$ — electron energy and momentum well-defined.

Nuclear Physics before 1930: nucleus = $n_p p + n_e e^-$.

Example: ${}^4\text{He} = 4p + 2e^-$, works well. However: ${}^{14}\text{N} = 14p + 7e^-$ is expected to be a fermion. However, it was experimentally known that ${}^{14}\text{N}$ was a boson!

There was also a problem with the magnetic moment of nuclei: $\mu_N, \mu_p \ll \mu_e$ ($\mu = eh/4mc$). How can the nuclear magnetic moment be so much smaller than the electron one if the nucleus contains electrons?

SOLUTION: Bound, nuclear electrons are very weird!

This can also be used to solve the continuous β -ray spectrum: energy need not be conserved in nuclear processes! (N. Bohr)

“... This would mean that the idea of **energy and its conservation fails** in dealing with processes involving the emission and capture of **nuclear electrons**. This does not sound improbable if we remember all that has been said about **peculiar properties of electrons in the nucleus**.” (G. Gamow, Nuclear Physics Textbook, 1931).

enter the neutrino...

1. 1930: Postulated by Pauli to (a) resolve the problem of continuous β -ray spectra, and (b) reconcile nuclear model with spin-statistics theorem. \Rightarrow
2. 1932: Chadwick discovers the neutron.
neutron \neq Pauli's neutron = neutrino (Fermi);
3. 1934: Fermi theory of Weak Interactions – current-current interaction

$$\mathcal{H} \sim G_F (\bar{p}\Gamma n) (\bar{e}\Gamma\nu_e), \quad \text{where } \Gamma = \{1, \gamma_5, \gamma_\mu, \gamma_\mu\gamma_5, \sigma_{\mu\nu}\}$$

Way to “see” neutrinos: $\bar{\nu}_e + p \rightarrow e^+ + n$. Prediction for the cross-section - too small to ever be observed...
4. 1935: (Yukawa postulates the existence of mesons (pions) as mediators of the nuclear (strong) force: $m_\pi \sim 100$ MeV.)
5. 1936/37: (“Meson” discovered in cosmic rays. Another long, tortuous story. Turns out to be the muon...)
6. 1947: (Marshak, Bethe postulate the 2 meson hypothesis ($\pi \rightarrow \mu$). Pion observed in cosmic rays.)

Dear Radioactive Ladies and Gentlemen,

I have come upon a desperate way out regarding the wrong statistics of the ^{14}N and ^6Li nuclei, as well as the continuous β -spectrum, in order to save the “alternation law” statistics and the energy law. To wit, the possibility that there could exist in the nucleus electrically neutral particles, which I shall call “neutrons,” and satisfy the exclusion principle... The mass of the neutrons should be of the same order of magnitude as the electron mass and in any case not larger than 0.01 times the proton mass. The continuous β -spectrum would then become understandable from the assumption that in β -decay a neutron is emitted along with the electron, in such a way that the sum of the energies of the neutron and the electron is constant... For the time being I dare not publish anything about this idea and address myself to you, dear radioactive ones, with the question how it would be with experimental proof of such a neutron, if it were to have the penetrating power equal to about ten times larger than a γ -ray.

I admit that my way out may not seem very probable *a priori* since one would probably have seen the neutrons a long time ago if they exist. But only the one who dares wins, and the seriousness of the situation concerning the continuous β -spectrum is illuminated by my honored predecessor, Mr Debye who recently said to me in Brussels: “Oh, it is best not to think about this at all, as with new taxes.” One must therefore discuss seriously every road to salvation. Thus, dear radioactive ones, examine and judge. Unfortunately, I cannot appear personally in Tübingen since a ball... in Zürich... makes my presence here indispensable... .

Your most humble servant, W. Pauli

Adapted summary of an English Translation to Pauli's letter dated December 4, 1930, from Ref. 3.

observing the unobservable:

1. 1956: “Discovery” of the neutrino (Reines and Cowan) in the Savannah River Nuclear Reactor site. \Rightarrow

$\bar{\nu}_e + p \rightarrow e^+ + n$. Measure positron ($e^+e^- \rightarrow \gamma s$) and neutron ($nN \rightarrow N^* \rightarrow N + \gamma s$) in delayed coincidence in order to get rid of backgrounds.

2. 1958: Neutrino Helicity Measured (Goldhaber et al.). Neutrinos are purely **left-handed**. Interact only weakly (Parity violated maximally).

$e^- + {}^{152}\text{Eu}(J=0) \rightarrow {}^{152}\text{Sm}^*(J=1) + \nu \rightarrow {}^{152}\text{Sm}(J=1) + \nu + \gamma$

3. 1962: The second neutrino: $\nu_\mu \neq \nu_e$ (Lederman, Steinberger, Schwartz at BNL). First neutrino beam.

$p + Z \rightarrow \pi^+ X \rightarrow \mu^+ \nu_\mu \Rightarrow$

$$\nu_\mu + Z \rightarrow \mu^- + Y \text{ (“always”)}$$

$$\nu_\mu + Z \rightarrow e^- + Y \text{ (“never”)}$$

4. 2001: ν_τ directly observed (DONUT experiment at FNAL). Same strategy: $\nu_\tau + Z \rightarrow \tau^- + Y$. (τ -leptons discovered in the 1970’s). \Rightarrow

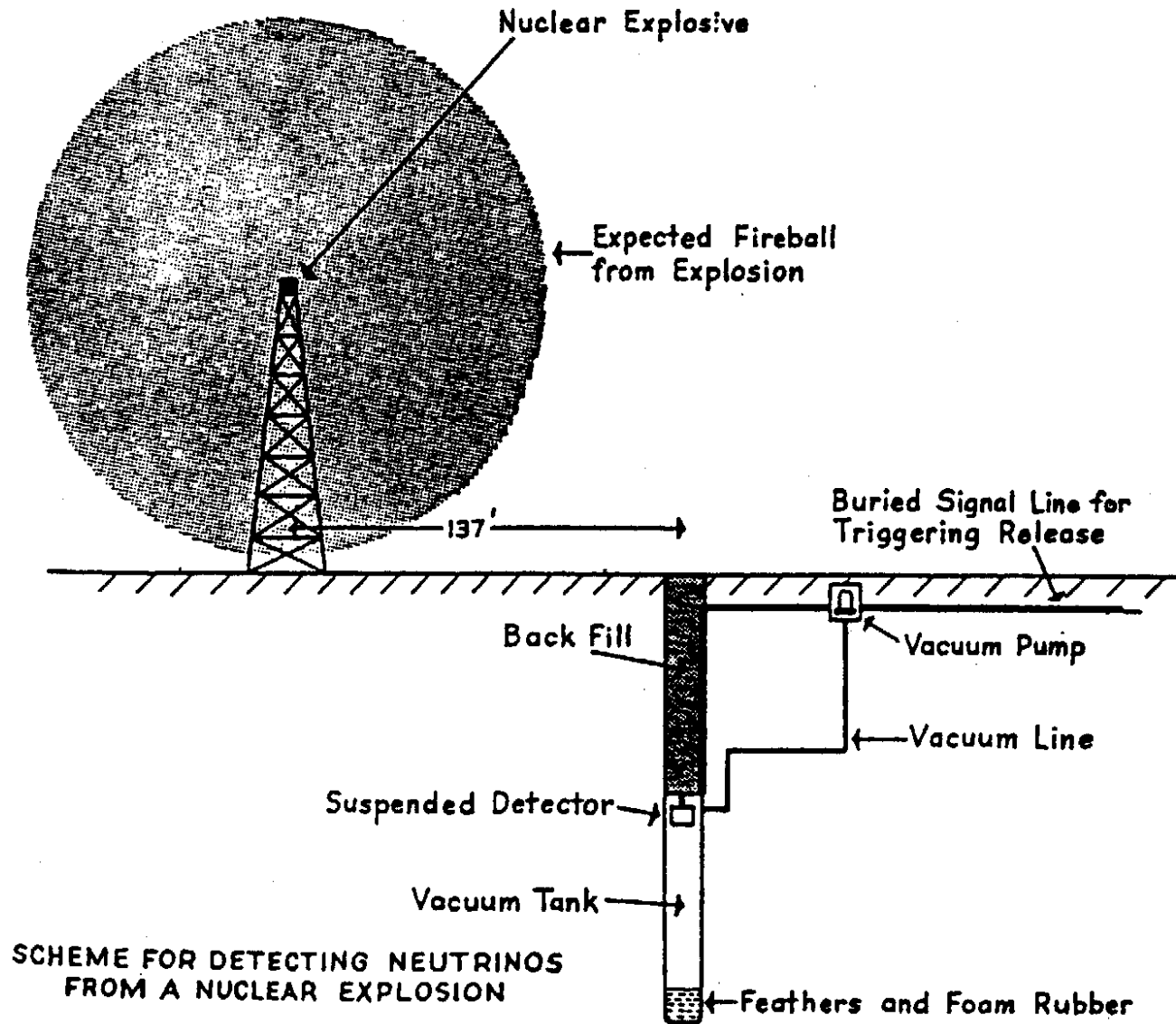


FIGURE 5.1 Scheme for detecting neutrinos from a nuclear explosion (Cowan, 1964).

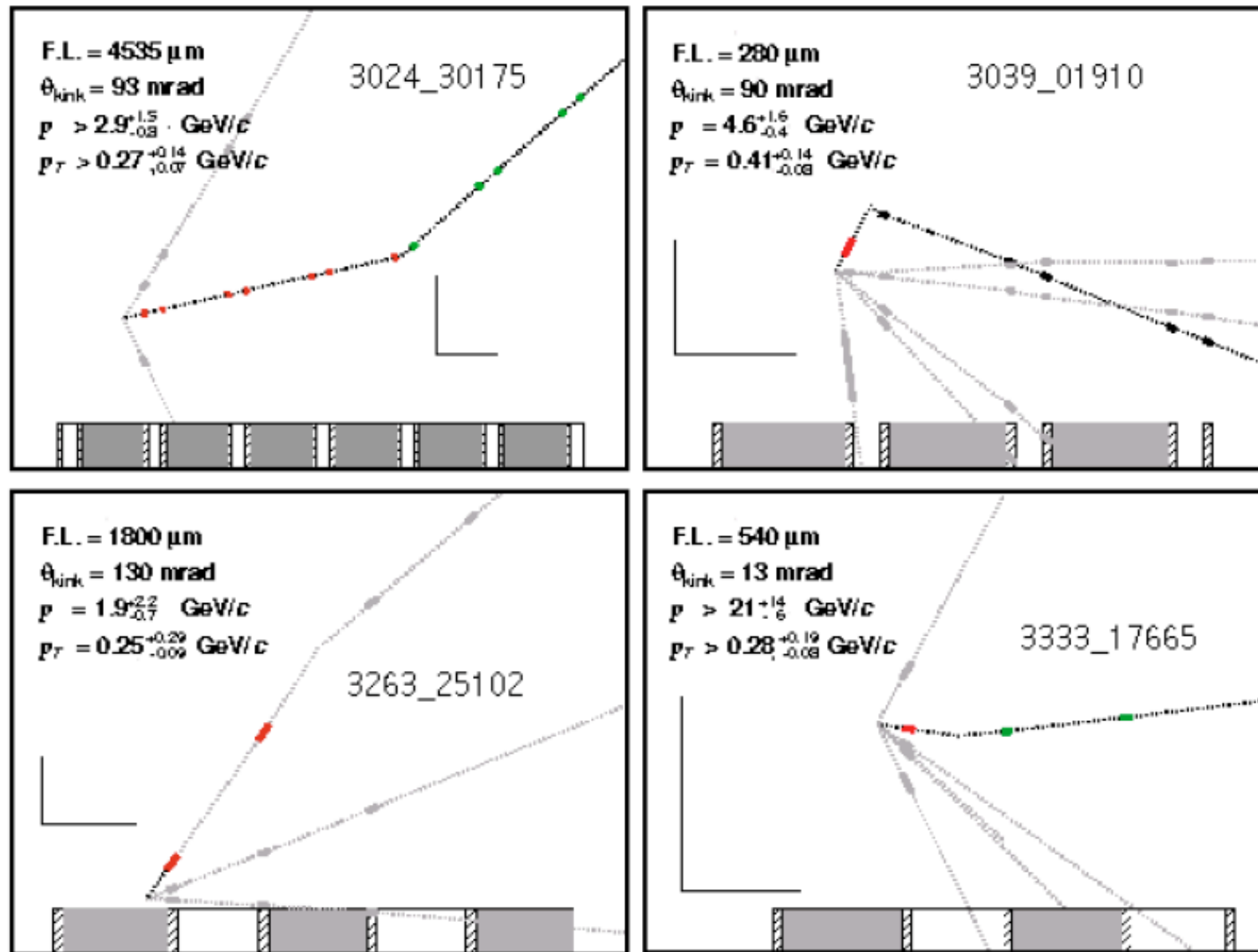


Figure 4-6: The four tau neutrino charged current events. The scale is given by the perpendicular lines (vertical: 0.1 mm, horizontal: 1 mm). The bar on the bottom shows the target material (solid: steel, hatched: emulsion, clear: plastic base).

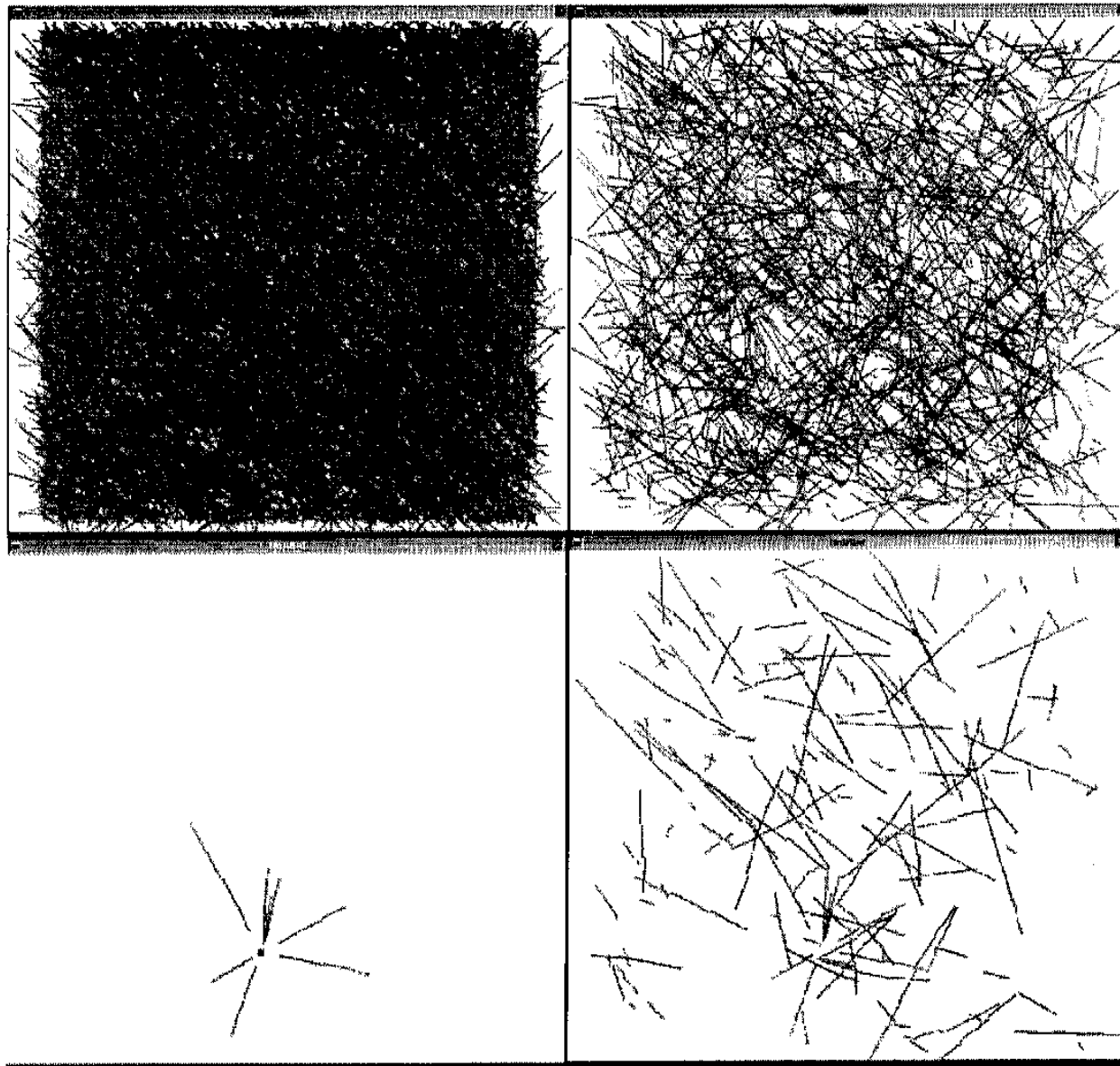
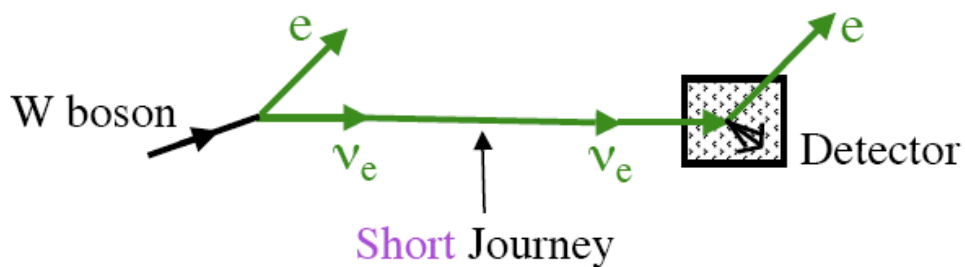


図 5.12: net scan 反応点探索の各段階 (左上から時計回り)。1) 読み込んだ全ての飛跡 ($5 \times 5\text{mm}^2$)、2) 測定領域を突き抜けている飛跡の排除、3) 低運動量の飛跡の排除、4) 一点 ($4\mu\text{m}$ 以内) 収束している飛跡

What we Knew of Neutrinos: End of the 20th Century



- come in three flavors (see figure);
- interact only via weak interactions (W^\pm, Z^0);
- have ZERO mass – helicity good quantum number;
- ν_L field describes 2 degrees of freedom:
 - left-handed state ν ,
 - right-handed state $\bar{\nu}$ (CPT conjugate);
- neutrinos carry lepton number:
 - $L(\nu) = +1$,
 - $L(\bar{\nu}) = -1$.

2– Neutrino Puzzles – 1960’s to 2000’s

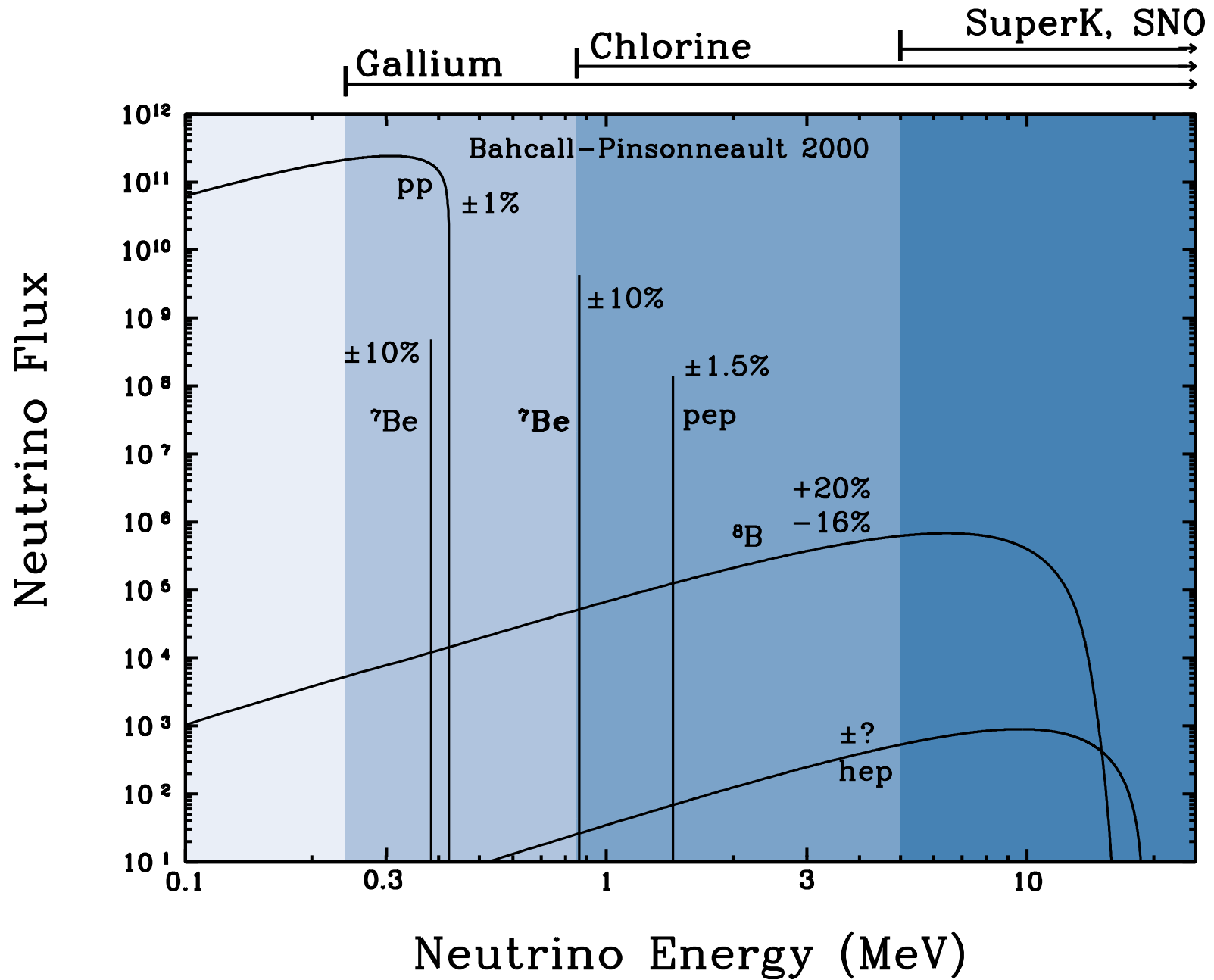
Long baseline neutrino experiments have revealed that **neutrinos change flavor** after propagating a finite distance, violating the definitions in the previous slide. The rate of change depends on the neutrino energy E_ν and the baseline L .

- $\nu_\mu \rightarrow \nu_\tau$ and $\bar{\nu}_\mu \rightarrow \bar{\nu}_\tau$ — atmospheric experiments [“indisputable”];
- $\nu_e \rightarrow \nu_{\mu,\tau}$ — solar experiments [“indisputable”];
- $\bar{\nu}_e \rightarrow \bar{\nu}_{\text{other}}$ — reactor neutrinos [“indisputable”];
- $\nu_\mu \rightarrow \nu_{\text{other}}$ — from accelerator experiments [“indisputable”].

Table 1. Nuclear reactions responsible for producing almost all of the Sun’s energy and the different “types” of solar neutrinos (nomenclature): *pp*-neutrinos, *pep*-neutrinos, *hep*-neutrinos, ${}^7\text{Be}$ -neutrinos, and ${}^8\text{B}$ -neutrinos. ‘Termination’ refers to the fraction of interacting protons that participate in the process.

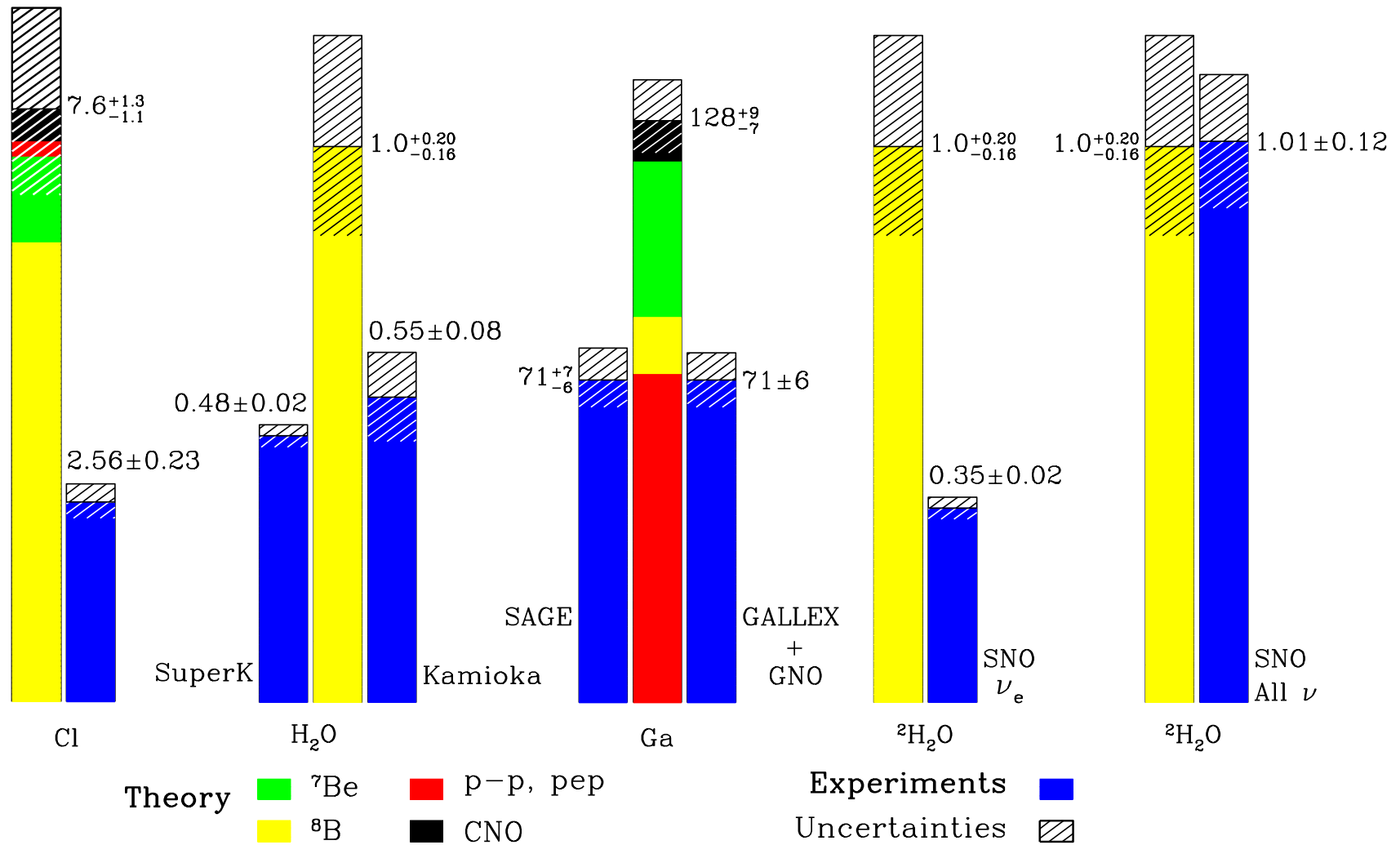
Reaction	Termination (%)	Neutrino Energy (MeV)	Nomenclature
$p + p \rightarrow {}^2\text{H} + e^+ + \nu_e$	99.96	< 0.423	<i>pp</i> -neutrinos
$p + e^- + p \rightarrow {}^2\text{H} + \nu_e$	0.044	1.445	<i>pep</i> -neutrinos
${}^2\text{H} + p \rightarrow {}^3\text{He} + \gamma$	100	–	–
${}^3\text{He} + {}^3\text{He} \rightarrow {}^4\text{He} + p + p$	85	–	–
${}^3\text{He} + {}^4\text{He} \rightarrow {}^7\text{Be} + \gamma$	15	–	–
${}^7\text{Be} + e^- \rightarrow {}^7\text{Li} + \nu_e$	15	0.863(90%) 0.386(10%)	${}^7\text{Be}$ -neutrinos
${}^7\text{Li} + p \rightarrow {}^4\text{He} + {}^4\text{He}$	–	–	–
${}^7\text{Be} + p \rightarrow {}^8\text{B} + \gamma$	0.02	–	–
${}^8\text{B} \rightarrow {}^8\text{Be}^* + e^+ + \nu_e$	–	< 15	${}^8\text{B}$ -neutrinos
${}^8\text{Be} \rightarrow {}^4\text{He} + {}^4\text{He}$	–	–	–
${}^3\text{He} + p \rightarrow {}^4\text{He} + e^+ + \nu_e$	0.00003	< 18.8	<i>hep</i> -neutrinos

Note: Adapted from Ref. 12. Please refer to Ref. 12 for a more detailed explanation.

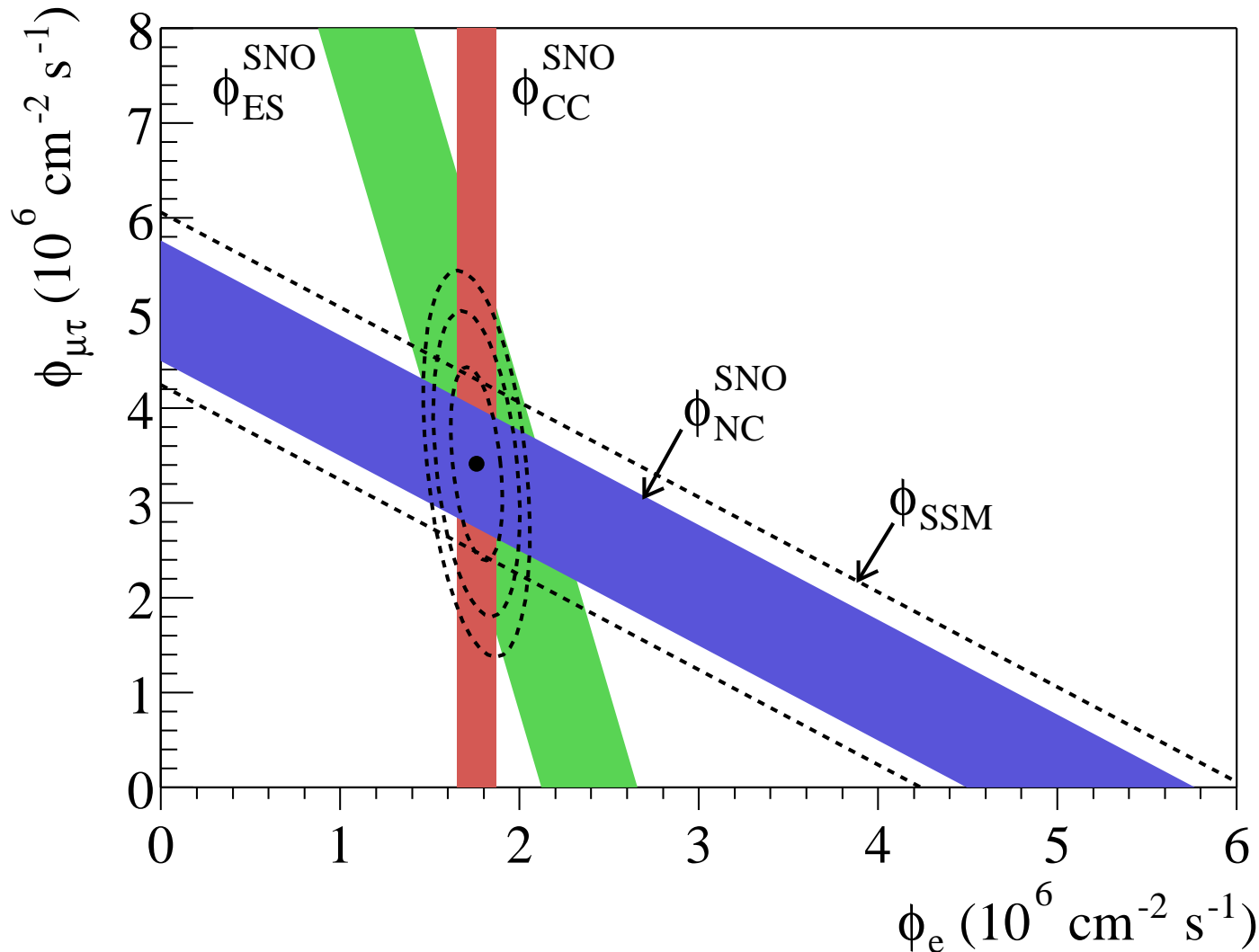


Total Rates: Standard Model vs. Experiment

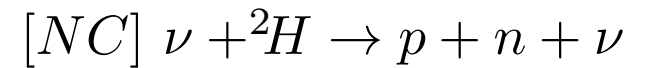
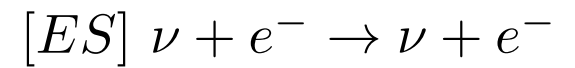
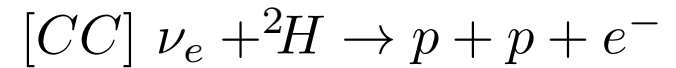
Bahcall-Pinsonneault 2000



The SNO Experiment: conclusive evidence for flavor change

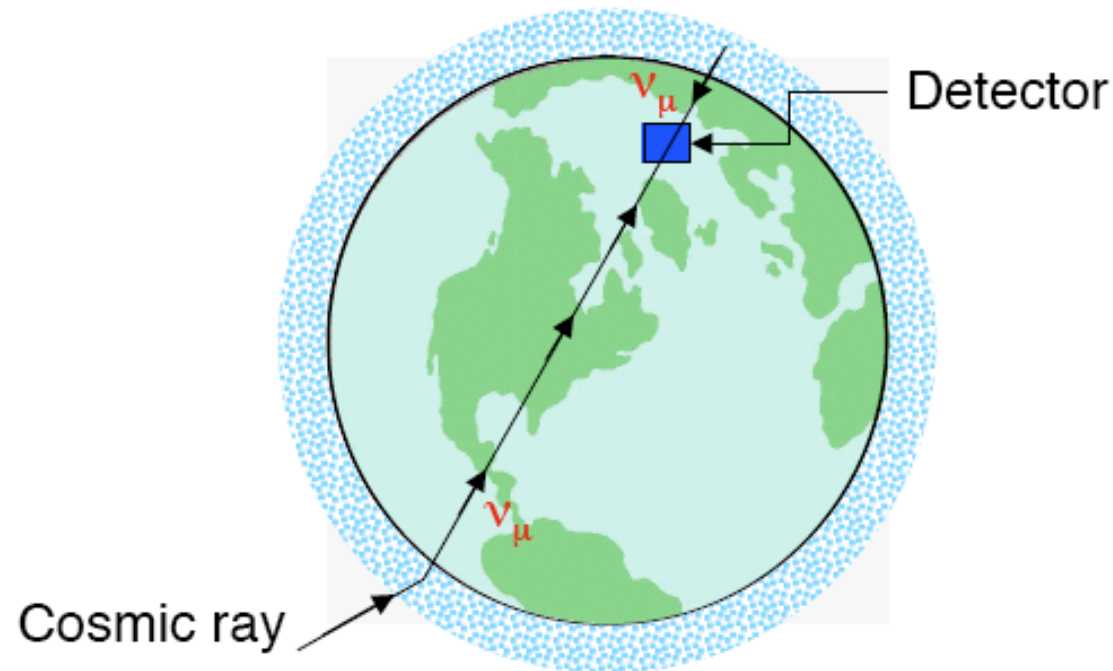


SNO Measures:



different reactions
sensitive to different
neutrino flavors.

Atmospheric Neutrinos



Isotropy of the $\gtrsim 2$ GeV cosmic rays + Gauss' Law + No ν_μ disappearance

$$\Rightarrow \frac{\phi_{\nu_\mu}(\text{Up})}{\phi_{\nu_\mu}(\text{Down})} = 1 .$$

But Super-Kamiokande finds for $E_\nu > 1.3$ GeV

$$\frac{\phi_{\nu_\mu}(\text{Up})}{\phi_{\nu_\mu}(\text{Down})} = 0.54 \pm 0.04 .$$

UP \neq DOWN – neutrinos can tell time! \rightarrow neutrinos have mass.

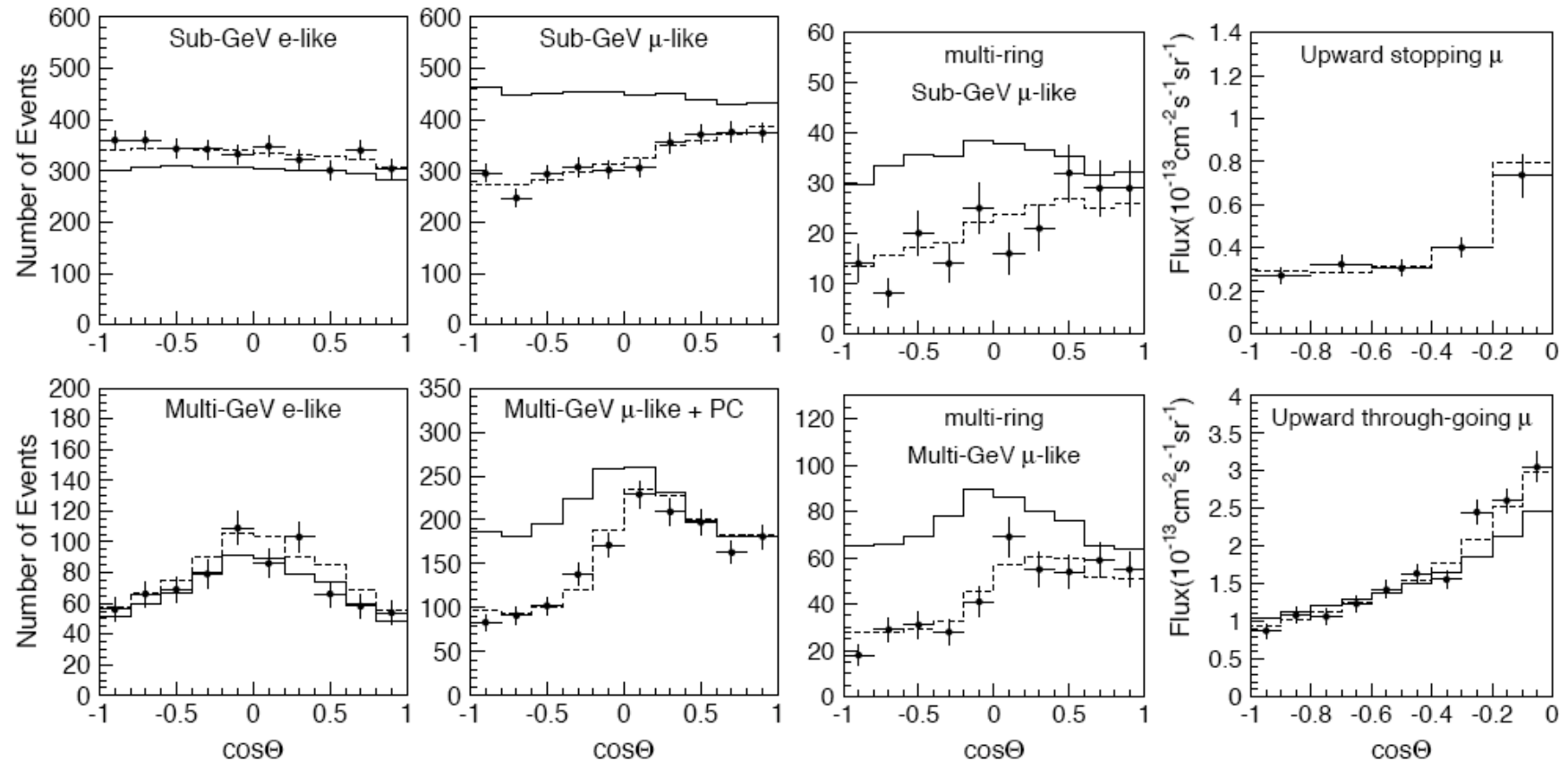


Figure 4. Zenith angle distribution for fully-contained single-ring e -like and μ -like events, multi-ring μ -like events, partially contained events and upward-going muons. The points show the data and the solid lines show the Monte Carlo events without neutrino oscillation. The dashed lines show the best-fit expectations for $\nu_\mu \leftrightarrow \nu_\tau$ oscillations. From M. Ishitsuka [Super-Kamiokande Collaboration], hep-ex/0406076.