Empirical Formula for the Maximum mass of { neutron stars with Relativistic mean-field theory

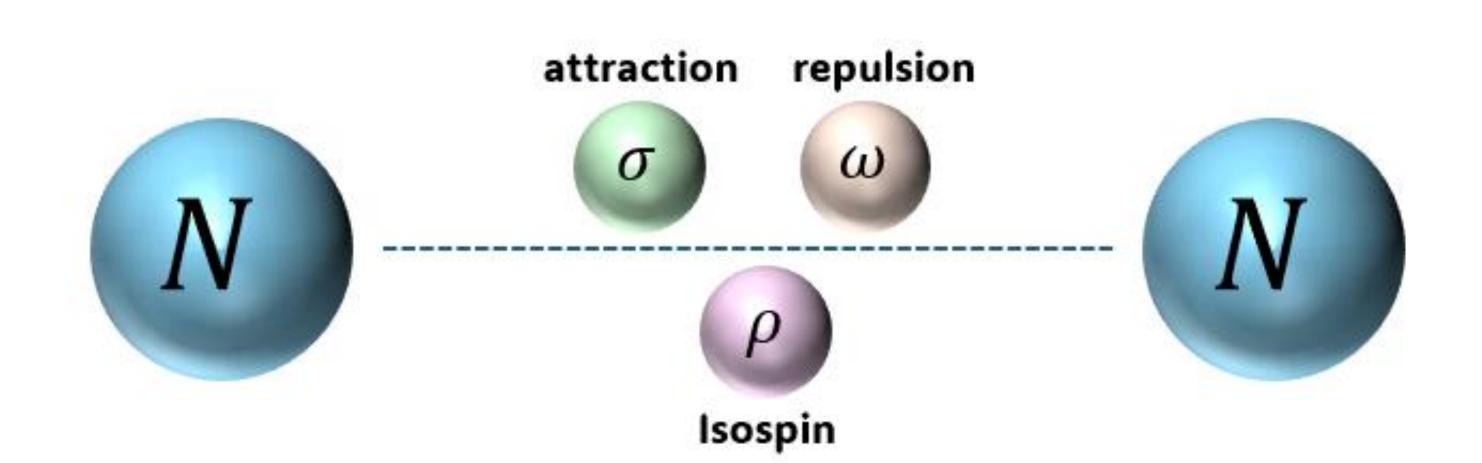
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Abstract

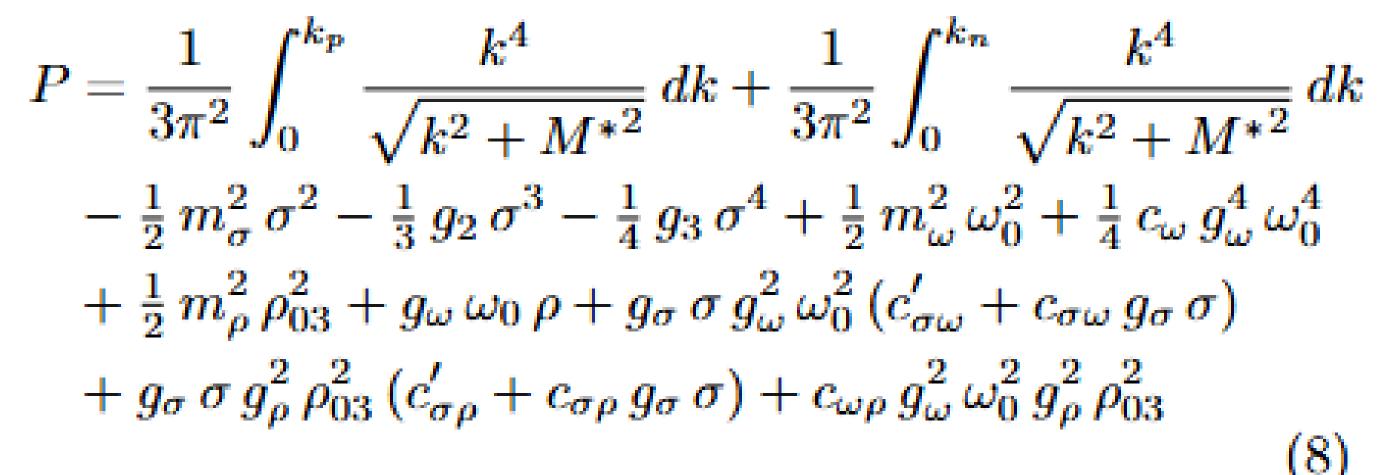
The empirical formula for the maximum mass of neutron stars is derived within the framework of the relativistic meanfield model. Observations of massive neutron stars heavier than $\sim 2\,M_{\odot}$ have ruled out soft equation of state and constrain nuclear interactions, which are valid in dense nuclear matter. Upon the request, numerous attempts have been made to refine the relativistic mean-field model, such as by including the delta meson. However, we find that the maximum mass of a neutron star predicted by the relativistic mean-field model can be primarily determined by the combination of the saturation density, the effective mass at saturation, and the vector meson self-coupling constant. While constraining the pure neutron matter equation of state using Chiral Effective Field Theory (ChEFT) at low densities, 250 parameter sets were generated to derive an empirical formula for the maximum mass of neutron stars and apply the formula with the present relativistic mean field models.

Relativistic Mean Field Theory

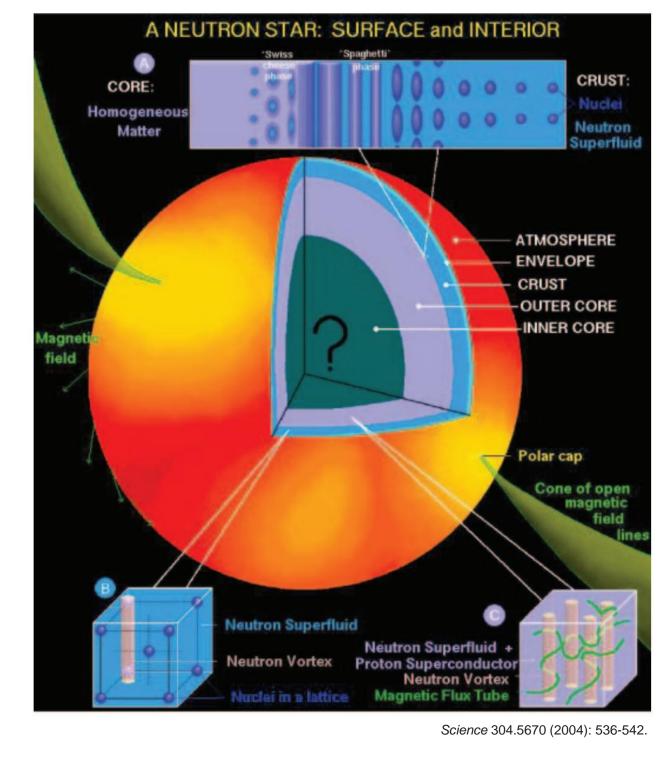


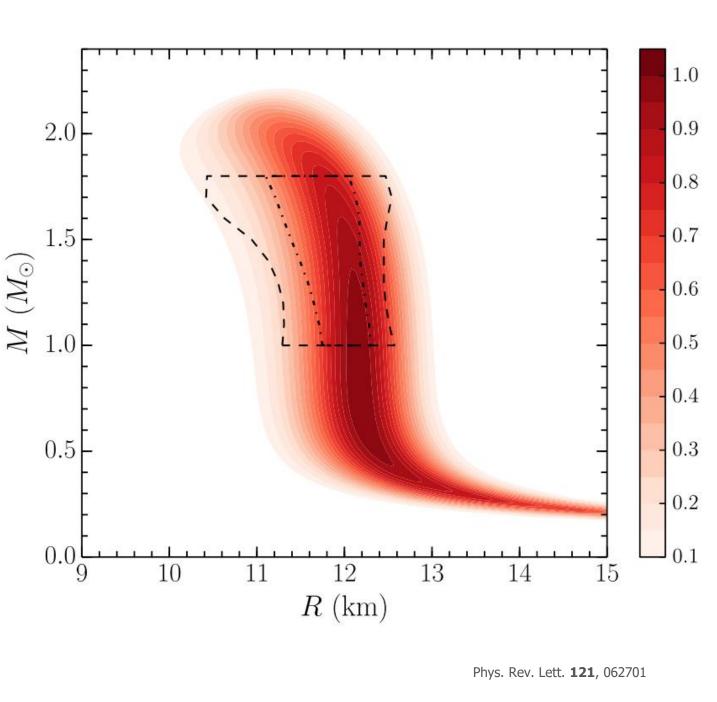
$$\mathcal{L} = \bar{\psi} (i\gamma^{\mu}\partial_{\mu} - M_{N})\psi + g_{\sigma} \,\sigma \,\bar{\psi}\psi - g_{\omega} \,\omega_{\mu} \,\bar{\psi}\gamma^{\mu}\psi
- g_{\rho} \,\rho_{\mu} \,\bar{\psi}\gamma^{\mu} \frac{\tau}{2}\psi + \frac{1}{2} \,\partial_{\mu}\sigma \,\partial^{\mu}\sigma - \frac{1}{2} \,m_{\sigma}^{2} \,\sigma^{2} - \frac{1}{3} \,g_{2} \,\sigma^{3}
- \frac{1}{4} \,g_{3} \,\sigma^{4} - \frac{1}{4} \,\omega_{\mu\nu} \,\omega^{\mu\nu} + \frac{1}{2} \,m_{\omega}^{2} \,\omega_{\mu} \,\omega^{\mu}
+ \frac{1}{4} \,c_{\omega} \,(g_{\omega}^{2} \,\omega_{\mu}\omega^{\mu})^{2} - \frac{1}{4} \,\rho_{\mu\nu} \,\rho^{\mu\nu} + \frac{1}{2} \,m_{\rho}^{2} \,\rho_{\mu} \,\rho^{\mu} \qquad (1)
+ g_{\sigma} \,\sigma \,g_{\omega}^{2} \,\omega_{\mu}\omega^{\mu} \,(c_{\sigma\omega}' + c_{\sigma\omega} \,g_{\sigma} \,\sigma)
+ g_{\sigma} \,\sigma \,g_{\rho}^{2} \,\rho_{\mu\rho} \,\rho^{\mu} \,(c_{\sigma\rho}' + c_{\sigma\rho} \,g_{\sigma} \,\sigma)
+ c_{\omega\rho} \,g_{\omega}^{2} \,\omega_{\mu}\omega^{\mu} \,g_{\rho}^{2} \,\rho_{\mu\rho}^{\mu} \,,$$

$$\varepsilon = \frac{1}{\pi^{2}} \int_{0}^{k_{p}} k^{2} \sqrt{k^{2} + M^{*2}} \,dk + \frac{1}{\pi^{2}} \int_{0}^{k_{n}} k^{2} \sqrt{k^{2} + M^{*2}} \,dk
+ \frac{1}{2} \,m_{\sigma}^{2} \,\sigma^{2} + \frac{1}{3} \,g_{2} \,\sigma^{3} + \frac{1}{4} \,g_{3} \,\sigma^{4}
- \frac{1}{2} \,m_{\omega}^{2} \,\omega_{0}^{2} - \frac{1}{4} \,c_{\omega} \,g_{\omega}^{4} \,\omega_{0}^{4} + g_{\omega} \,\omega_{0} \,\rho
- \frac{1}{2} \,m_{\rho}^{2} \,\rho_{03}^{2} + \frac{1}{2} \,g_{\rho} \,\rho_{03} \,\rho_{3} - g_{\sigma} \,\sigma \,g_{\omega}^{2} \,\omega_{0}^{2} \,(c_{\sigma\omega}' + c_{\sigma\omega} \,g_{\sigma} \,\sigma)
- g_{\sigma} \,\sigma \,g_{\rho}^{2} \,\rho_{03}^{2} \,(c_{\sigma\rho}' + c_{\sigma\rho} \,g_{\sigma} \,\sigma) - c_{\omega\rho} \,g_{\omega}^{2} \,\omega_{0}^{2} \,g_{\rho}^{2} \,\rho_{03}^{2}, \tag{7}$$



Neutron star

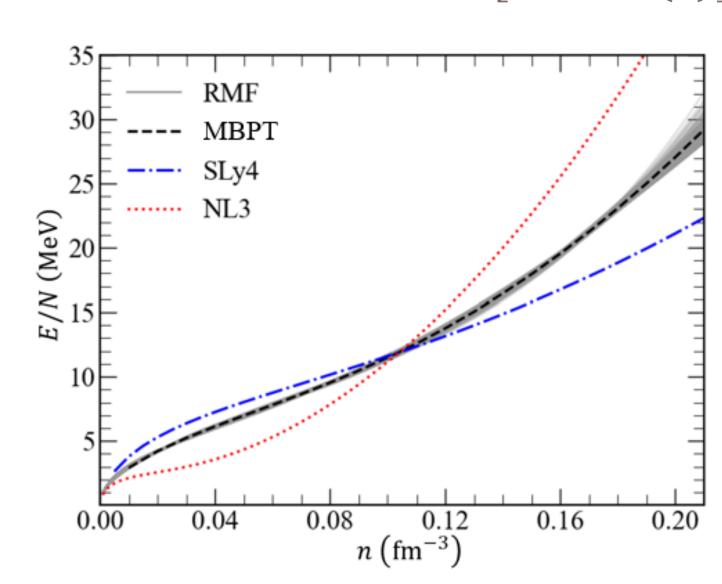


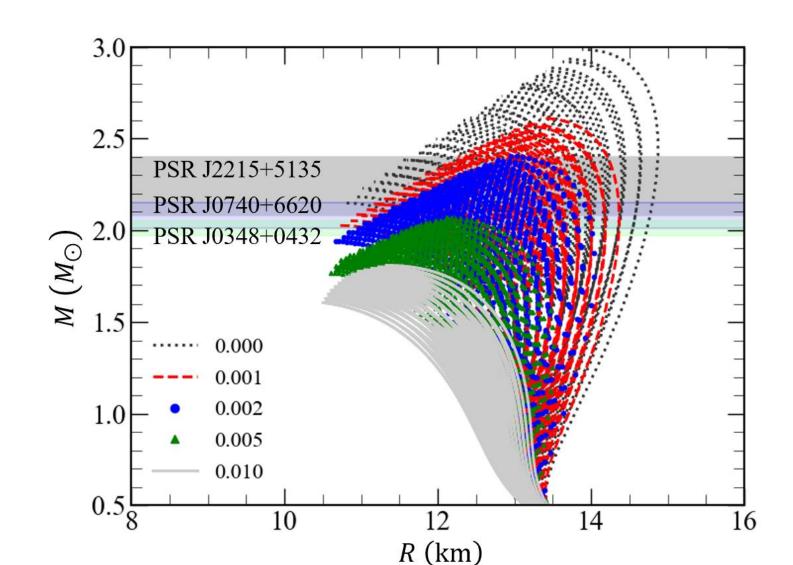


Results

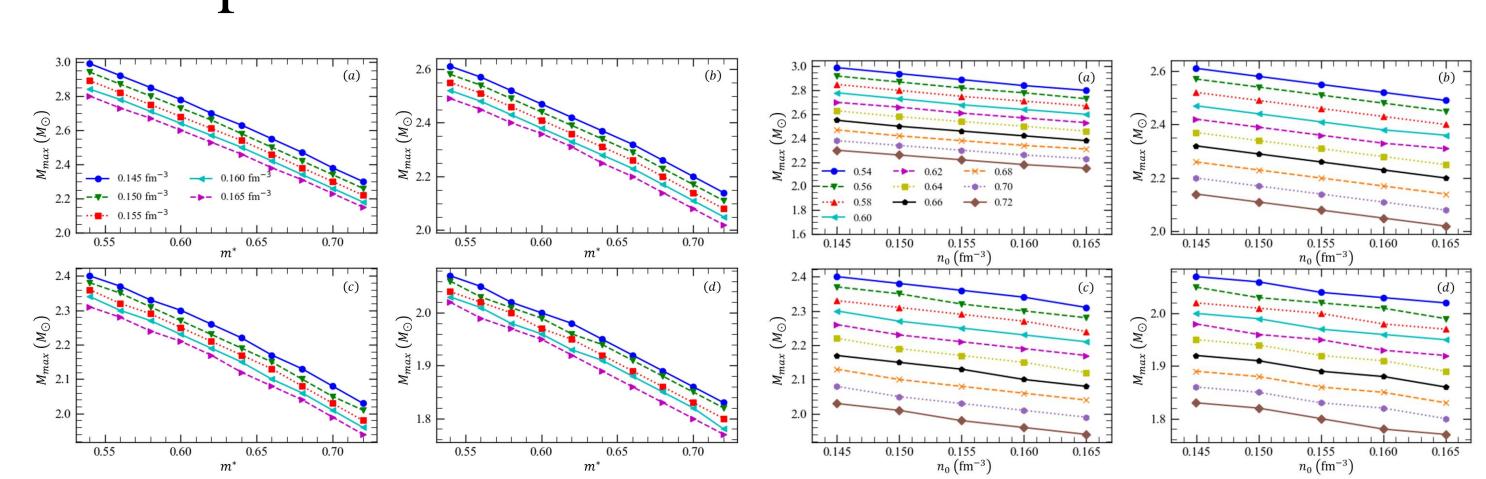
■TOV neutron stars

$$\frac{dP}{dr} = -\frac{[P(r) + \epsilon(r)][M(r) + 4\pi r^3 P(r)]}{r[r - 2M(r)]} \qquad \frac{dM}{dr} = 4\pi r^2 \epsilon(r)$$

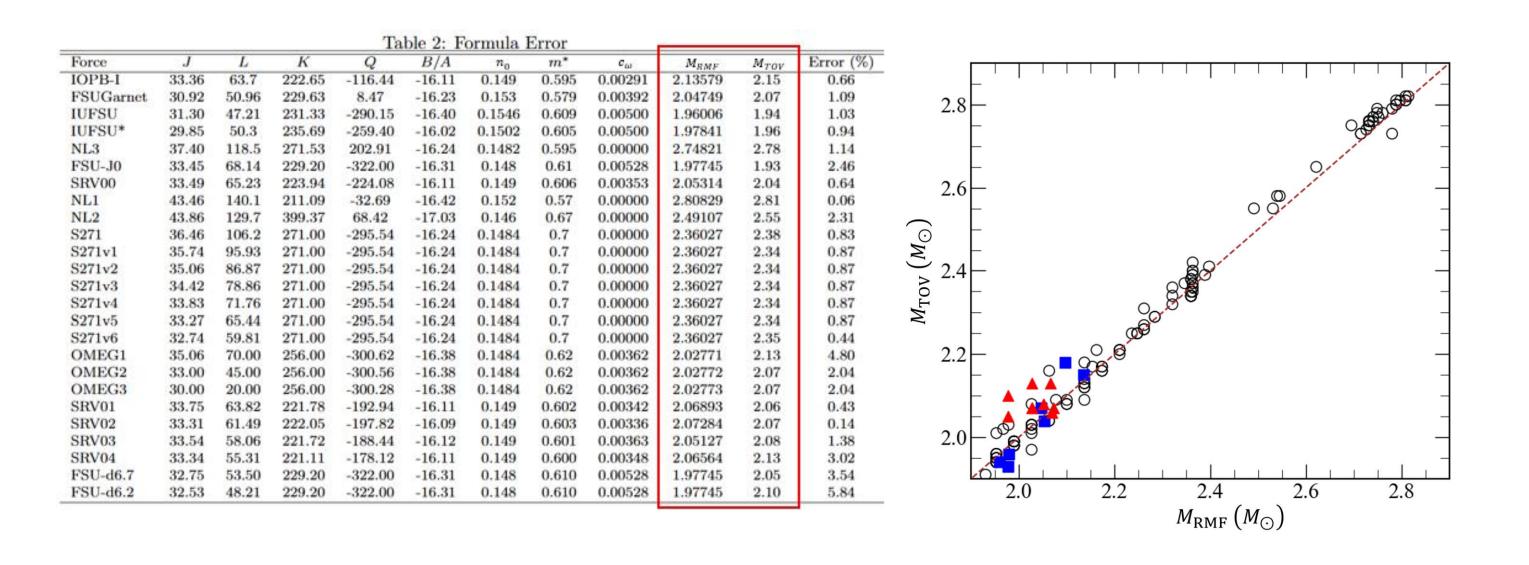


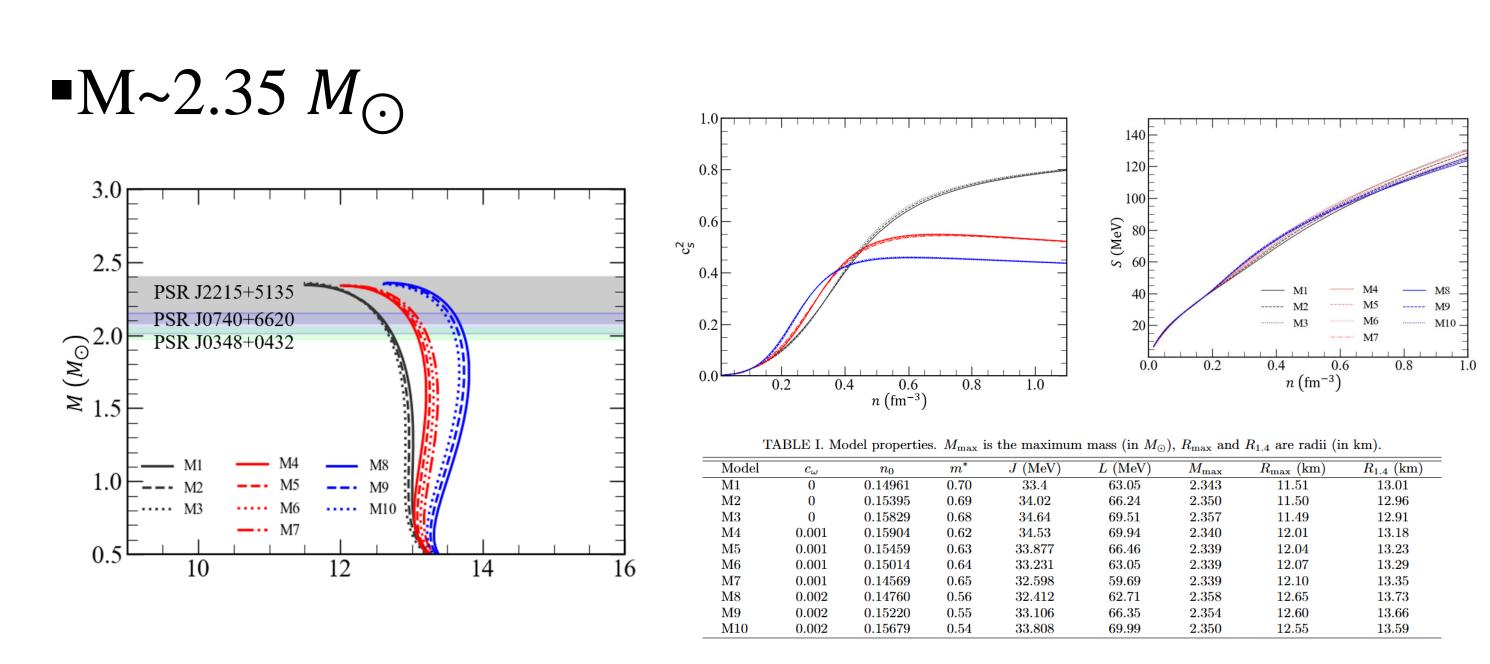


Empirical Formula



$$M_{max}(n_0, m^*, c_{\omega}) = f(c_{\omega})n_0 + g(c_{\omega})m^* + h(c_{\omega})$$





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