Particle Physics Models for DM-DR interactions

Pyungwon Ko (KIAS)

Based on P.Ko,Y.Tang; **1608.01083(PLB)**
1609.02307(PLB)
(*P.Ko,N.Nagata,Y.Tang;* **arXiv:1706.05605**)

DSU 2017, IBS-CTPU
Daejon, Korea (2017)

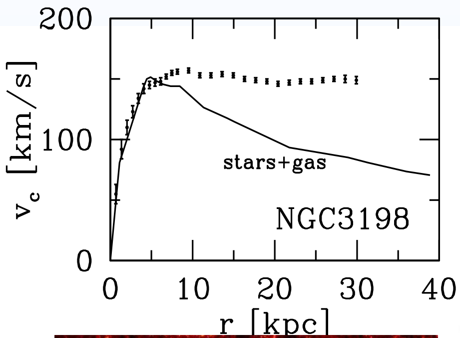
Outline

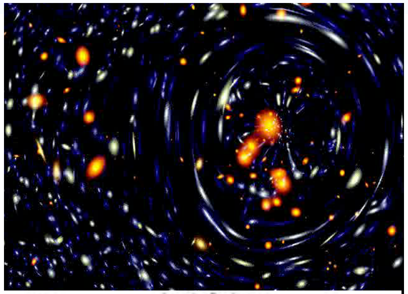
- Introduction & Motivation
  - Dark Matter evidence
  - Hubble constant and structure growth
- DM with dark gauge symmetries
- Interacting Dark Matter&Dark Radiation
  - U(1) dark photon
  - Residual Yang-Mills Dark Matter
- Summary

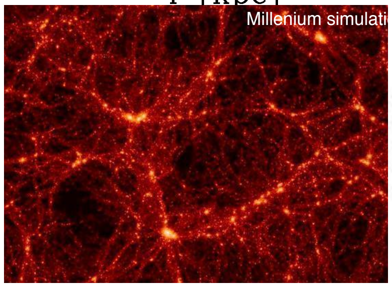
Only Higgs (~SM) and Nothing Else at the LHC & SM based on local gauge principle works very well !

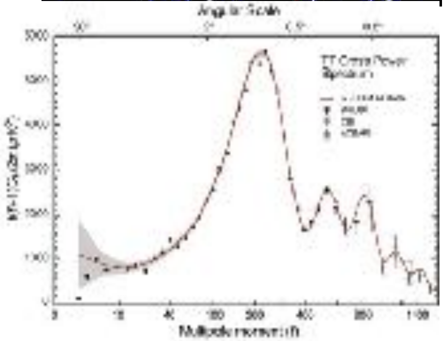
Dark Matter Evidence

- Rotation Curves of Galaxies
- Gravitational Lensing
- Large Scale Structure
- CMB anisotropies, …



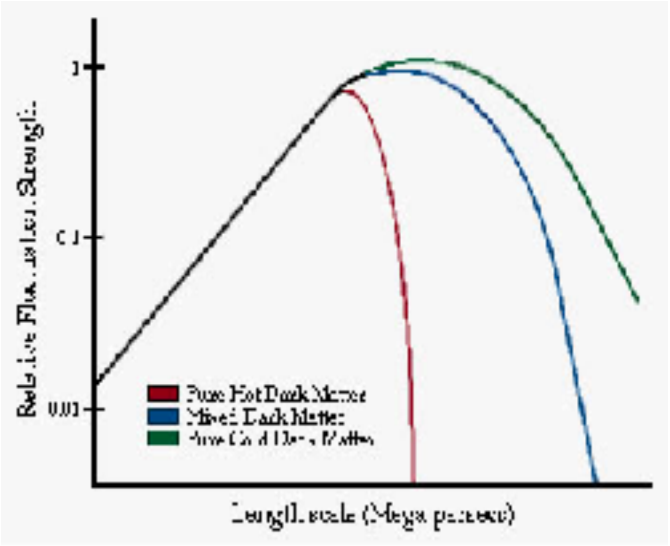






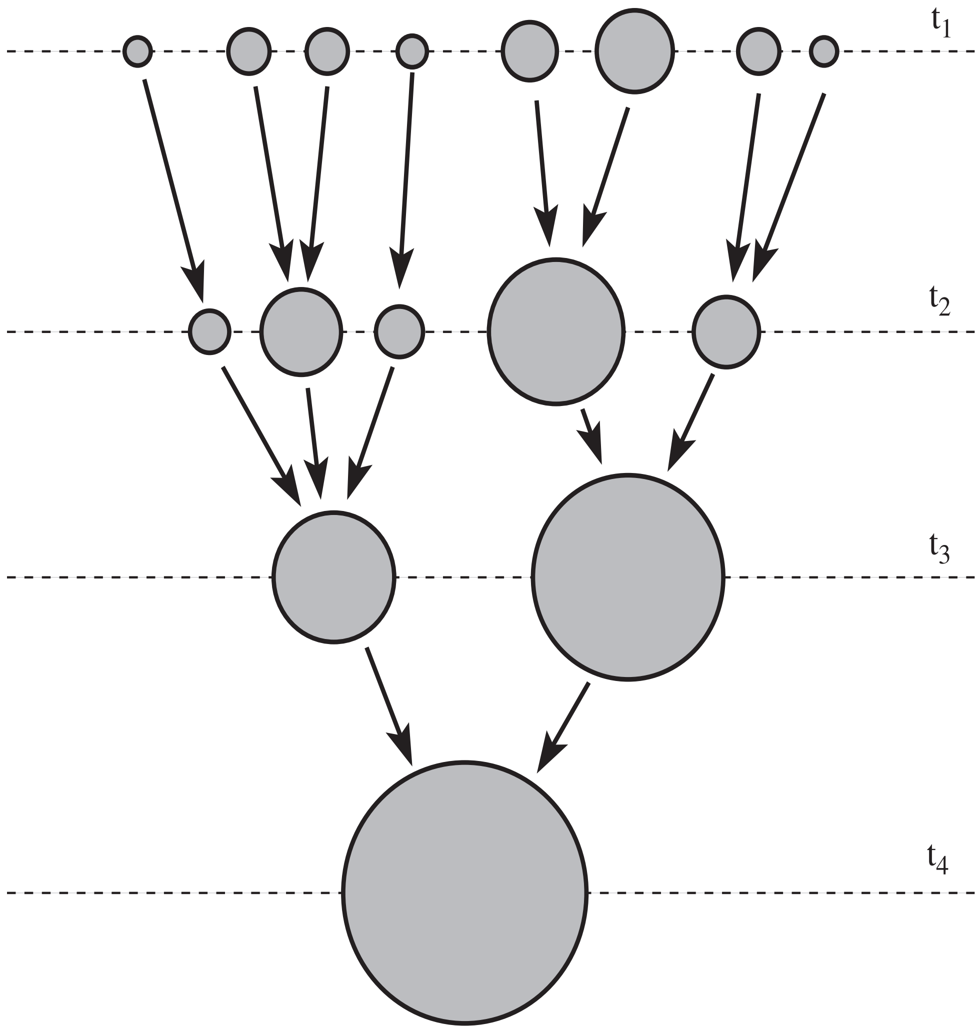
All *confirmed* *evidence comes from gravitational interaction*

CDM: negligible velocity, WIMP
WDM: keV sterile neutrino
HDM: active neutrino



Merger History of Dark Halo

- Standard picture
- DM halo grow hierarchically
- Small scale structures form first
- then merge into larger halo



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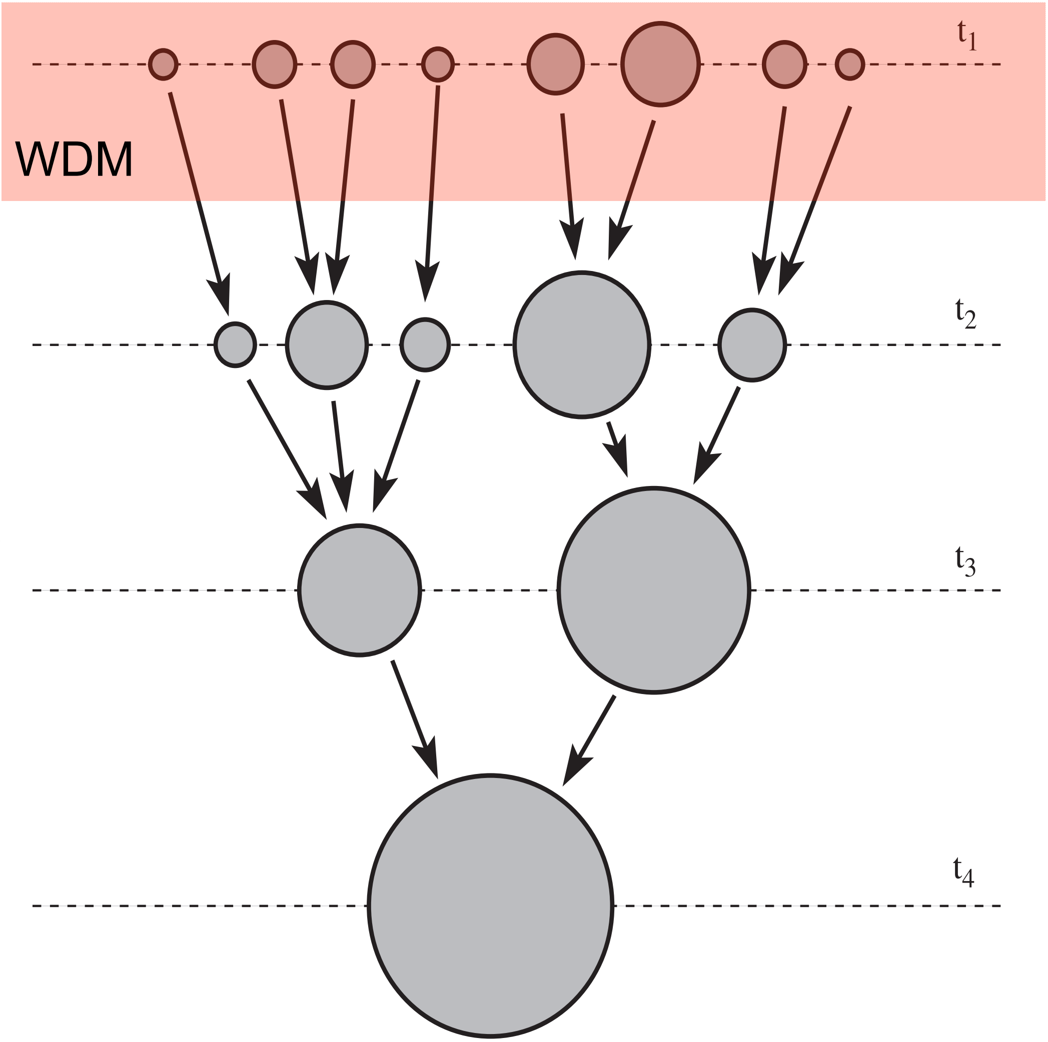

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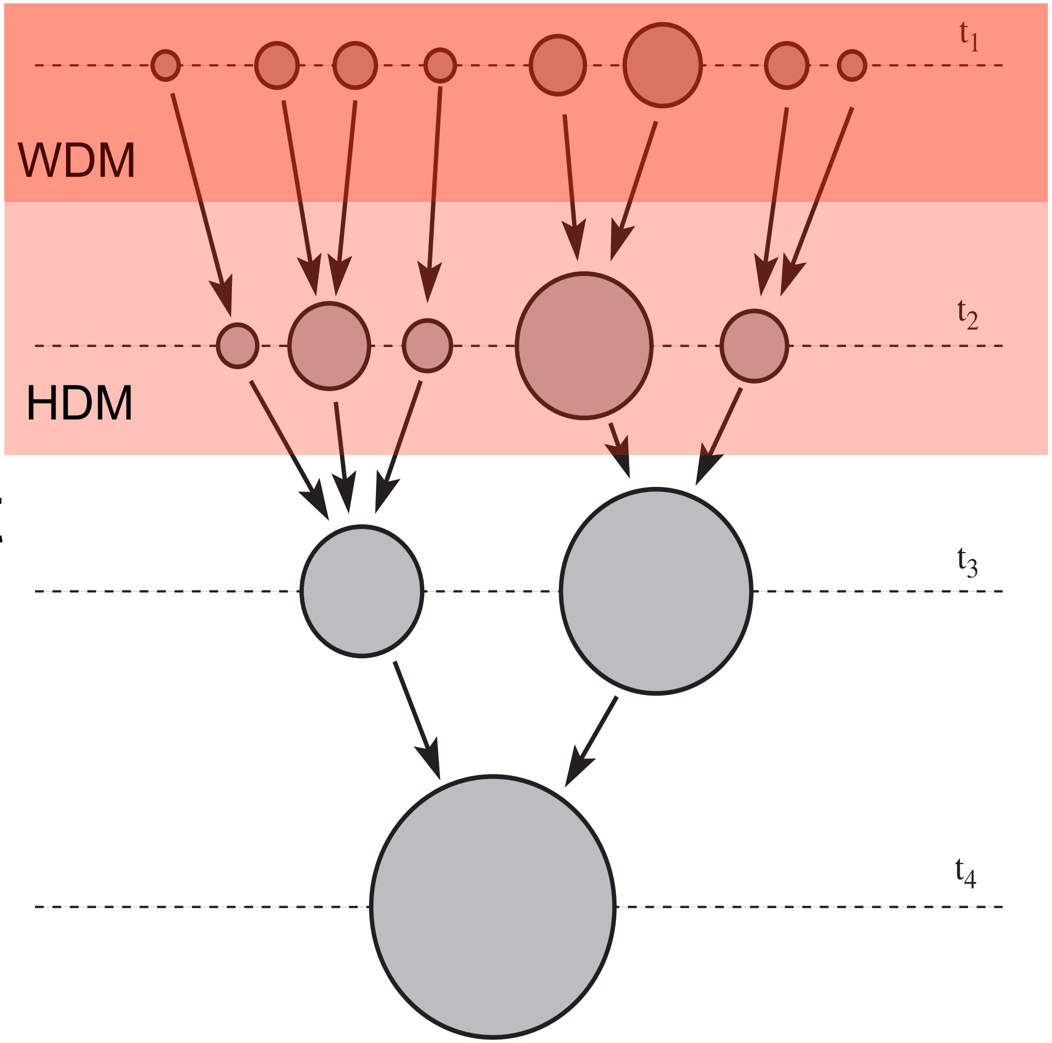
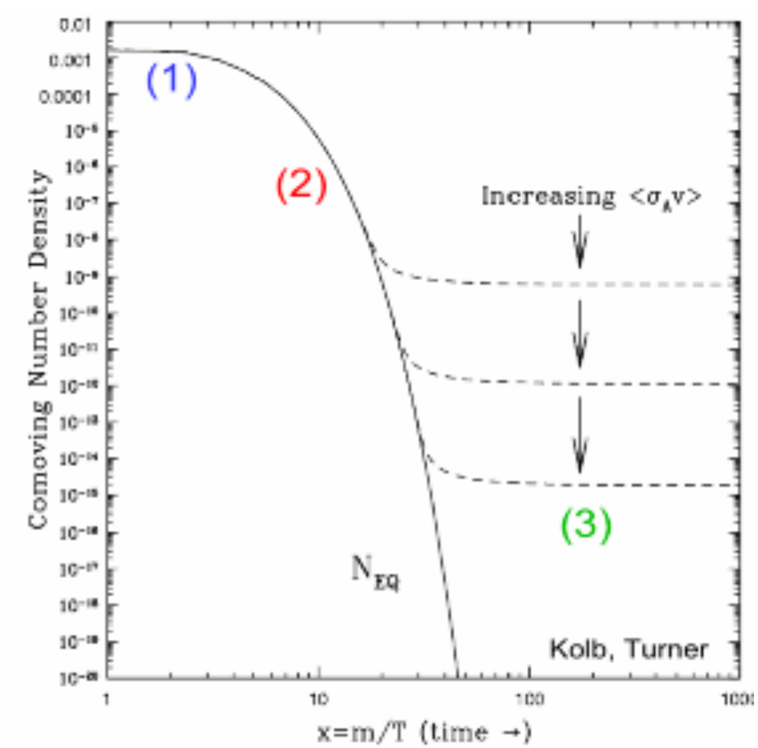
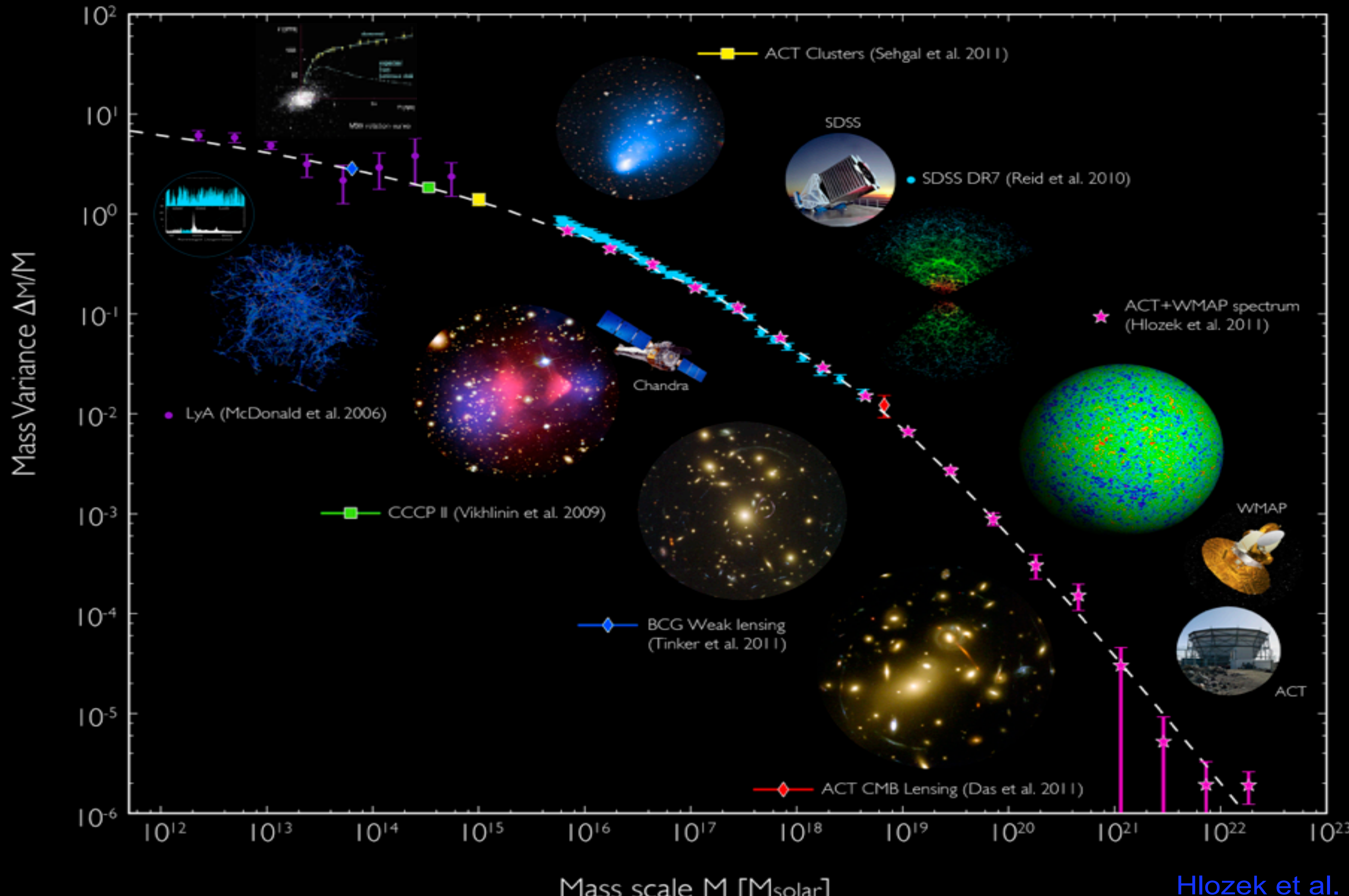


Weakly Interacting Massive Particle

- Mass around ~100GeV
- Coupling ~ 0.5
- Correct relic abundance Ω~0.3
- Thermal History
  - Equilibrium XX<>ff
  - Equilibrium XX >ff
  - Freeze-out
- Cold Dark Matter (CDM)



ΛCDM: successful on large scales



Hlozek et al. 2012

Interacting Dark Matter

Why Interacting DM ?

- Theoretically interesting
  - Atomic DM, Mirror DM, Composite DM
  - Eventually, all DM is *interacting* in some way, the question is how strongly?
  - Self-Interacting DM $\frac{\sigma}{M_X} \sim \mathrm{cm}^2/\mathrm{g} \sim \mathrm{barn}/\mathrm{GeV}$
- Possible new testable signatures
  - *CMB, LSS, BBN*
  - Other astrophysical effects,…
- Solution of CDM controversies
  - *Cusp-vs-Core, Too-big-to-fail, missing satellite,…*
  - H_0, σ_8? 2-3σ, systematic uncertainty

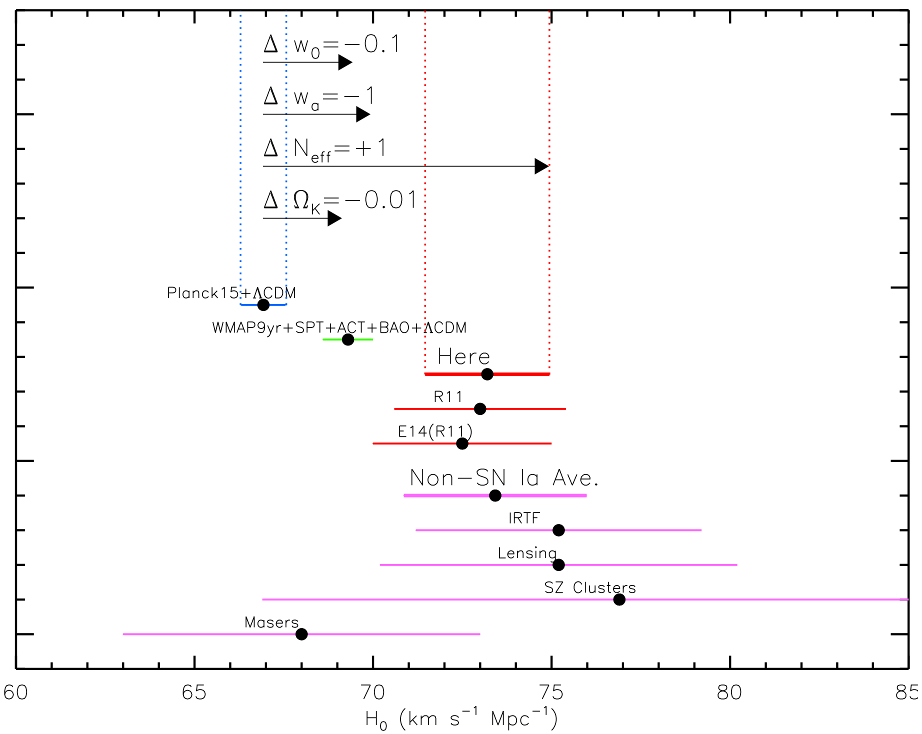
Tension in Hubble Constant?

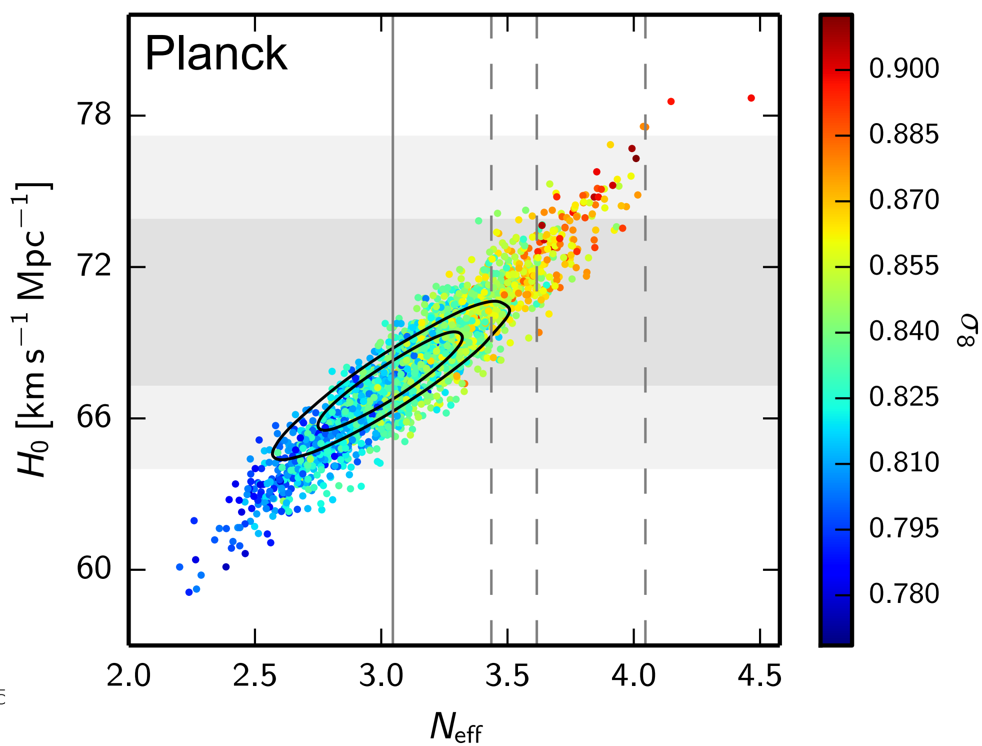
- Hubble Constant H_0 defined as the present value of

$$H \equiv \frac{1}{a}\frac{da}{dt} = \frac{\sqrt{\rho_r + \rho_m + \rho_\Lambda}}{M_p}$$

- `Planck(2015) gives` 67.8 ± 0.9 km s^{-1}Mpc^{-1}
- `HST(2016)    gives` 73.24 ± 1.74 km s^{-1}Mpc^{-1}

RIESS ET AL.





Tension in σ_8 ?

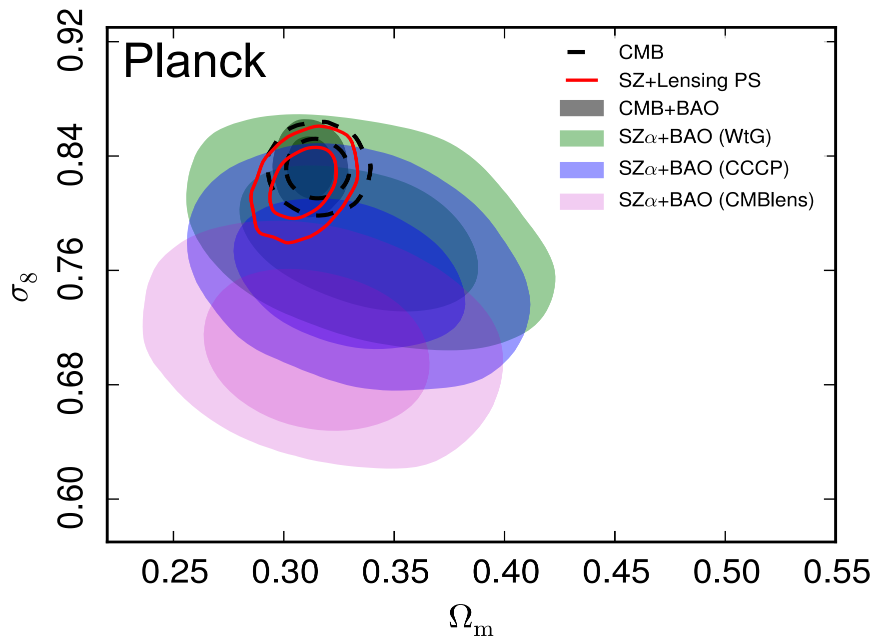
- Variance of perturbation field → collapsed objects

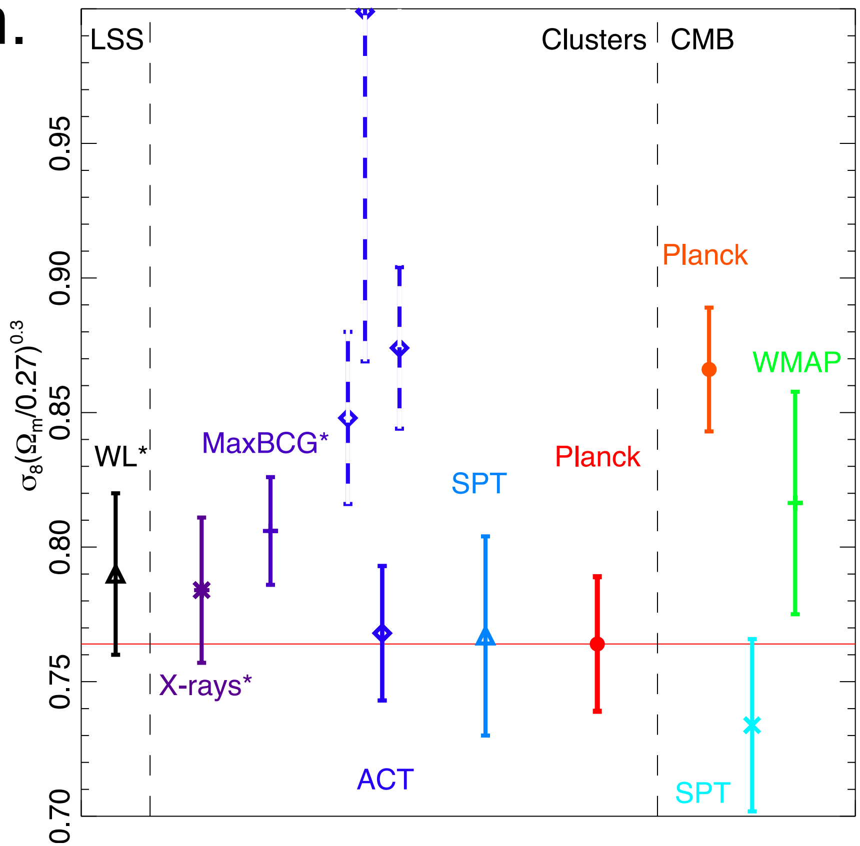
$$\sigma^2(R) = \frac{1}{2\pi^2}\int W_R^2(k)P(k)k^2dk,$$

- where the filter function $W_R(k) = \frac{3}{(kR)^3}\left[\sin(kR) - kR\cos(kR)\right],$

  $P(k)$ is matter power spectrum.

- $\sigma_8 \equiv \sigma(8h^{-1}\mathrm{Mpc})$





Tension in σ_8 ?

Planck2015, Sunyaev–Zeldovich cluster counts

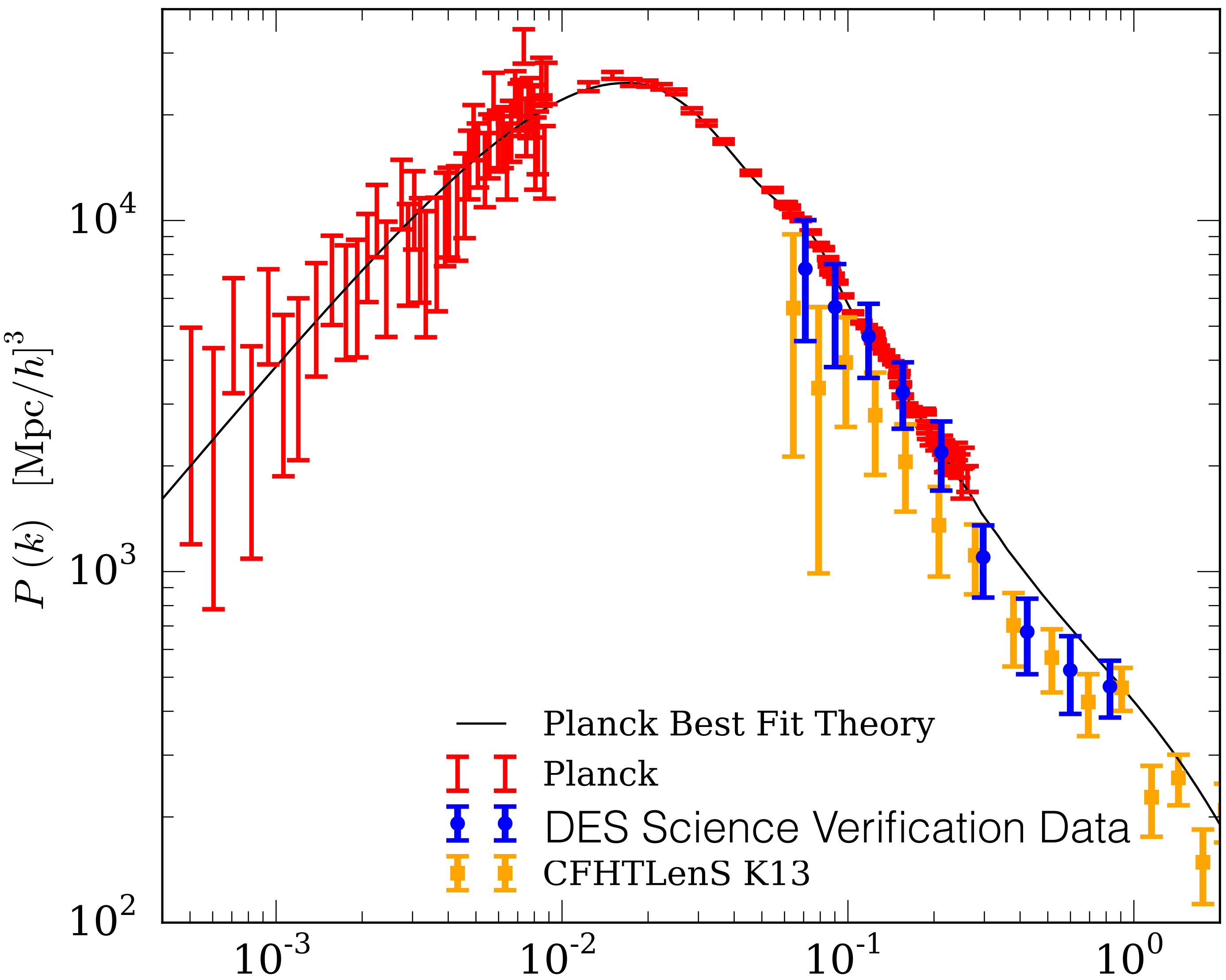
Data	$\sigma_8 \left(\frac{\Omega_m}{0.31}\right)^{0.3}$	Ω_m	σ_8
WtG + BAO + BBN	0.806 ± 0.032	0.34 ± 0.03	0.78 ± 0.03
CCCP + BAO + BBN **[Baseline]**	0.774 ± 0.034	0.33 ± 0.03	0.76 ± 0.03
CMBlens + BAO + BBN	0.723 ± 0.038	0.32 ± 0.03	0.71 ± 0.03
CCCP + H_0 + BBN	0.772 ± 0.034	0.31 ± 0.04	0.78 ± 0.04

Planck2015, Primary CMB

Parameter	[1] *Planck* TT+lowP	[2] *Planck* TE+lowP	[3] *Planck* EE+lowP	[4] *Planck* TT,TE,EE+lowP
$\Omega_b h^2$	0.02222 ± 0.00023	0.02228 ± 0.00025	0.0240 ± 0.0013	0.02225 ± 0.00016
$\Omega_c h^2$	0.1197 ± 0.0022	0.1187 ± 0.0021	$0.1150^{+0.0048}_{-0.0055}$	0.1198 ± 0.0015
$100\theta_{MC}$	1.04085 ± 0.00047	1.04094 ± 0.00051	1.03988 ± 0.00094	1.04077 ± 0.00032
τ	0.078 ± 0.019	0.053 ± 0.019	$0.059^{+0.022}_{-0.019}$	0.079 ± 0.017
$\ln(10^{10} A_s)$	3.089 ± 0.036	3.031 ± 0.041	$3.066^{+0.046}_{-0.041}$	3.094 ± 0.034
n_s	0.9655 ± 0.0062	0.965 ± 0.012	0.973 ± 0.016	0.9645 ± 0.0049
H_0	67.31 ± 0.96	67.73 ± 0.92	70.2 ± 3.0	67.27 ± 0.66
Ω_m	0.315 ± 0.013	0.300 ± 0.012	$0.286^{+0.027}_{-0.038}$	0.3156 ± 0.0091
σ_8	0.829 ± 0.014	0.802 ± 0.018	0.796 ± 0.024	0.831 ± 0.013
$10^9 A_s e^{-2\tau}$	1.880 ± 0.014	1.865 ± 0.019	1.907 ± 0.027	1.882 ± 0.012

Matter Power Spectrum

DES astroph/150705552

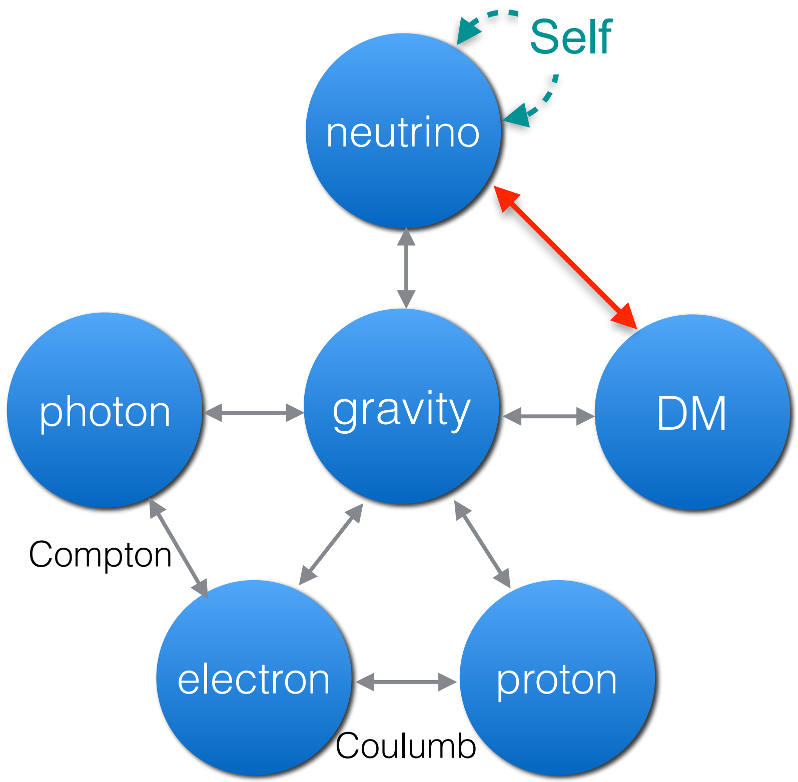


Interacting DM-DR

- Since all components are connected by Einstein's equation

$$R_{\mu\nu} - \frac{1}{2}Rg_{\mu\nu} + \Lambda g_{\mu\nu} = \frac{8\pi G}{c^4}T_{\mu\nu}$$

- first-order perturbation of Boltzmann equation
  - anisotropy in CMB
  - matter power spectrum for LSS
- (Self-)Interaction sometimes also matters



Interacting Radiation

- free-streaming

$$\dot{\delta}_\nu = -\frac{4}{3}\theta_\nu + 4\dot{\phi} ,$$

$$\dot{\theta}_\nu = k^2\left(\frac{1}{4}\delta_\nu - \sigma_\nu\right) + k^2\psi ,$$

$$\dot{F}_{\nu l} = \frac{k}{2l+1}\left[lF_{\nu(l-1)} - (l+1)F_{\nu(l+1)}\right] ,$$

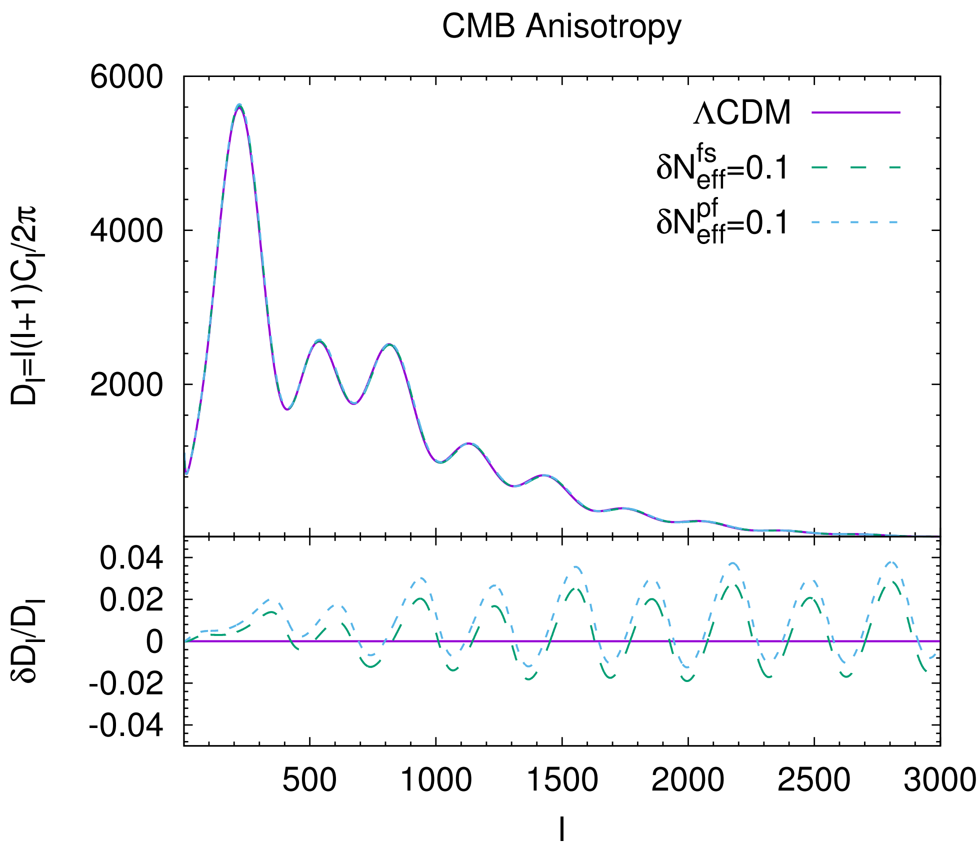
- perfect fluid $\Gamma \gg \mathcal{H}$

$$\dot{\delta}_\nu = -\frac{4}{3}\theta_\nu + 4\dot{\phi} ,$$

$$\dot{\theta}_\nu = k^2\left(\frac{1}{4}\delta_\nu - \sigma_\nu\right) + k^2\psi ,$$

$\sigma_\nu = 0$

Y.Tang, arXiv:1603.00165(PLB)



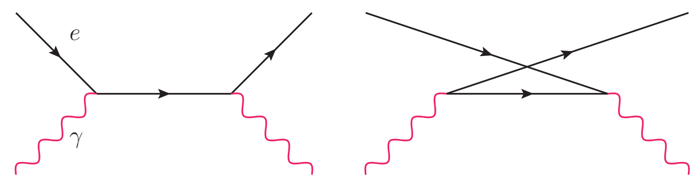
Neutrinos as perfect fluid excluded,
Audren **et al 1412.5948**

Relation to Particle Physics

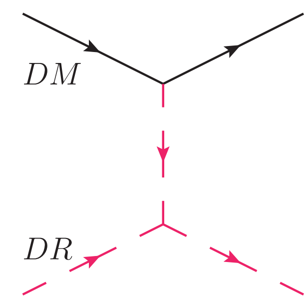
- The precise form of the scattering term, $\langle \sigma c \rangle$, is fully determined by the underlying microscopic or particle physics model, for example

IR behaviour

  - electron-photon, $\langle \sigma c \rangle \sim 1/m^2$
    Thomson scattering -> CMB, BAO



  - DM-radiation with massive mediator, $\langle \sigma c \rangle \sim T^2/m^4$
    Boehm *et al*(astro-ph/0410591,1309.7588)



  - non-Abelian radiation, $\langle \sigma c \rangle \sim 1/T^2$
    Schmaltz et al(2015), 1507.04351,1505.03542
  - (pseudo-)scalar radiation, $\langle \sigma c \rangle \sim 1/T^2$, μ^2/T^4, T^2/μ^4
    Y.Tang,1603.00165(PLB)

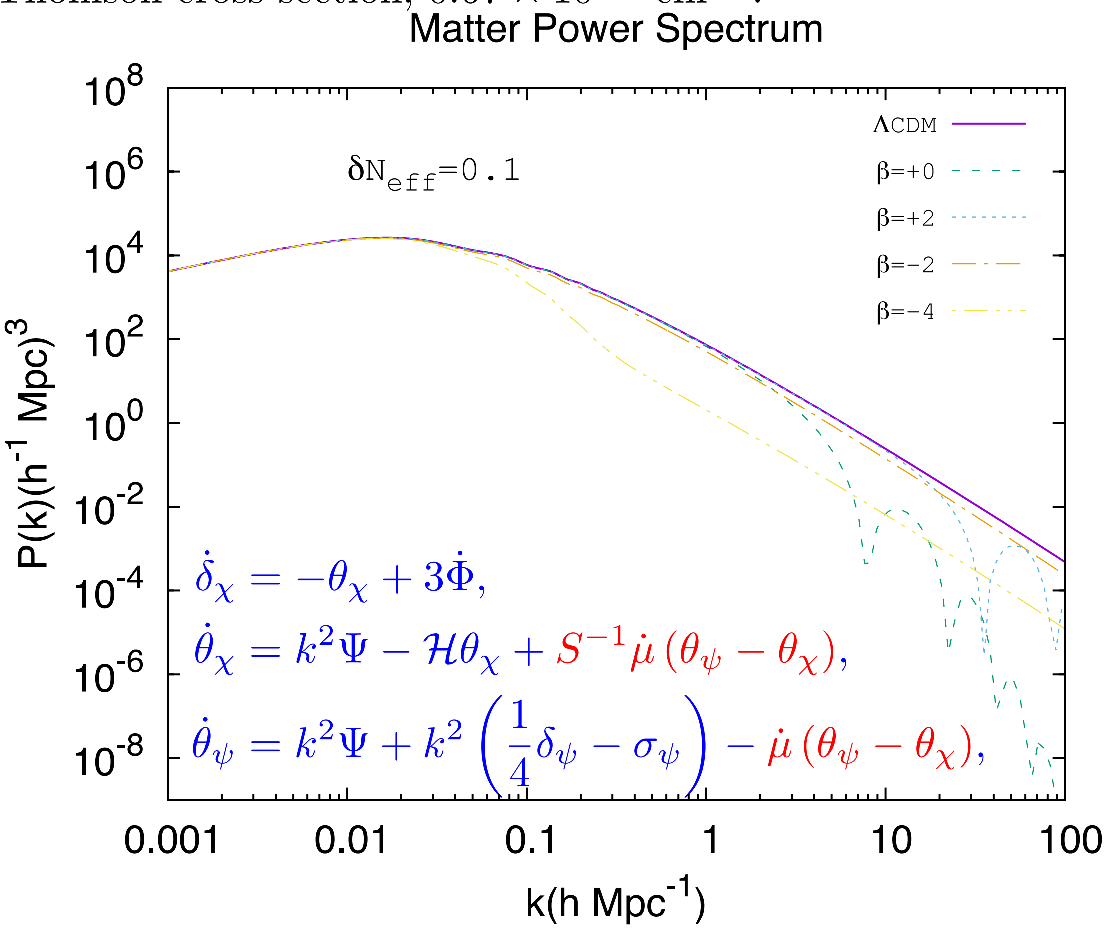
Effects on LSS

Parametrize the cross section ratio

Y.Tang,1603.00165(PLB)

$$u_0 \equiv \left[\frac{\sigma_{\chi\psi}}{\sigma_{\rm Th}}\right]\left[\frac{100{\rm GeV}}{m_\chi}\right], u_\beta(T) = u_0\left(\frac{T}{T_0}\right)^\beta,$$

where $\sigma_{\rm Th}$ is the Thomson cross section, $0.67 \times 10^{-24}{\rm cm}^{-2}$.



$$\dot\delta_\chi = -\theta_\chi + 3\dot\Phi,$$

$$\dot\theta_\chi = k^2\Psi - \mathcal{H}\theta_\chi + S^{-1}\dot\mu\left(\theta_\psi - \theta_\chi\right),$$

$$\dot\theta_\psi = k^2\Psi + k^2\left(\frac{1}{4}\delta_\psi - \sigma_\psi\right) - \dot\mu\left(\theta_\psi - \theta_\chi\right),$$

Why dark gauge sym ?

Questions about DM

- Electric Charge/Color neutral
- How many DM species are there ?
- Their masses and spins ?
- Are they absolutely stable or very long lived ?
- How do they interact with themselves and with the SM particles ?
- Where do their masses come from ? Another (Dark) Higgs mechanism ? Dynamical SB ?
- In order to answer these questions, we must find DM in particle physics experiments (direct/indirect detections, collider searches, etc.) and study their properties

DM phenomenology often requires

- New force mediators (scalar, vector, ….) in order to solve some puzzles in the standard collision less CDM paradigm
- Extra particles in the dark sector (excited DM, dark radiation, force mediators, etc.) often used for phenomenological reasons
- Any good organizing principles for these extra particles ?
- Answer : Dark gauge symmetry (dark gauge boson/dark Higgs appear naturally, their dynamics is completely fixed by gauge principle)

What is going on in the SM ?

- SM based on Poincare + local gauge symmetry within 4-dim QFT : extremely successful and provides qualitative answers to light neutrino masses, non-observation of proton decay (Lepton # and baryon # : accidental symmetry of the renormalizable SM, and broken only by higher dim operators)
- Electron is stable, because electric charge is conserved and electron is the lightest particle with nonzero electric charge
- Proton is long lived because B-violation in SM comes from dim-6 operator

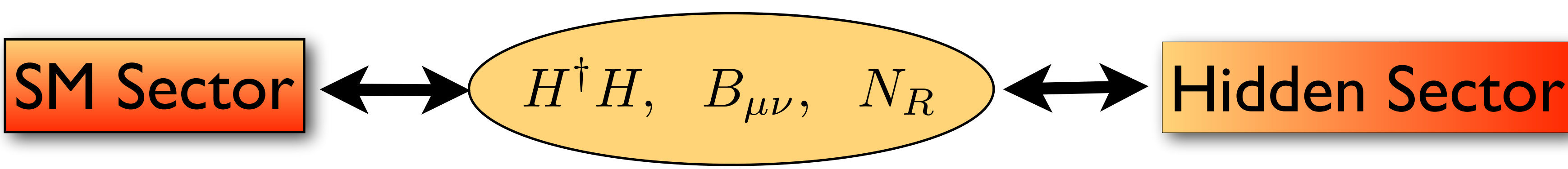
DM with dark gauge symmetries

- DM : either absolutely stable or long lived (could be due to local gauge symmetry or some accidental symmetry) and both can be accommodated by local dark gauge symmetries
- Global sym could be broken by gravity, and may not be good enough for DM stability/longevity
- The only issue is the mass scales of DM, dark gauge bosons/dark Higgs, and their gauge/ Yukawa couplings, all of which are unknown yet
- DM phenomenology can be very rich, if these new particles are not too heavy

Singlet Portal

- If there is a hidden (dark) sector with its own dark gauge symmetry and DM is thermal, then we need a portal to it
- There are only three unique gauge singlets in the SM + RH neutrinos

Baek, Ko, Park, arXiv:1303.4280, JHEP

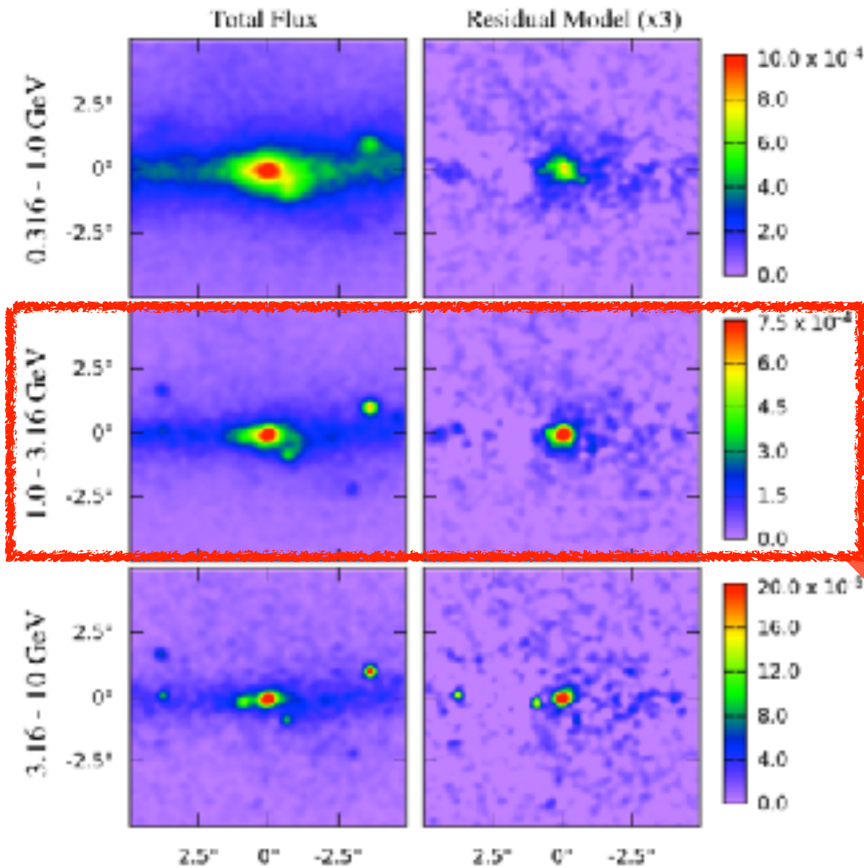


$$\boxed{N_R \leftrightarrow \widetilde{H} l_L}$$

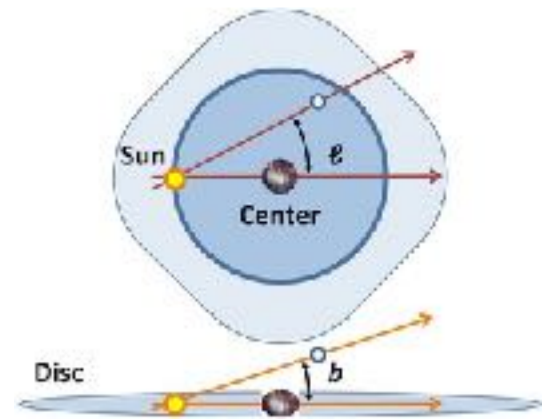
$$\boxed{e.g.\ \phi_X^\dagger \phi_X, X_{\mu\nu}, \psi_X^\dagger \phi_X}$$

Example: Fermi-LAT γ-ray excess

- Gamma-ray excess in the direction of GC



[1402.6703, T. Daylan et.al.]



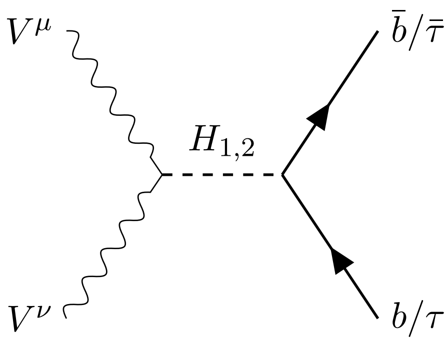
$$\mathrm{GC}: b \sim l \lesssim 0.1^\circ$$

extended

GeV scale excess!

GC gamma ray in VDM

[1404.5257, P. Ko, WIP & Y. Tang] JCAP (2014)
(Also Celine Boehm et al. 1404.4977, PRD)



H2 : 125 GeV Higgs
H1 : present in VDM
with dark gauge sym

Figure 2. Dominant s channel $b + \bar{b}$ (and $\tau + \bar{\tau}$) production

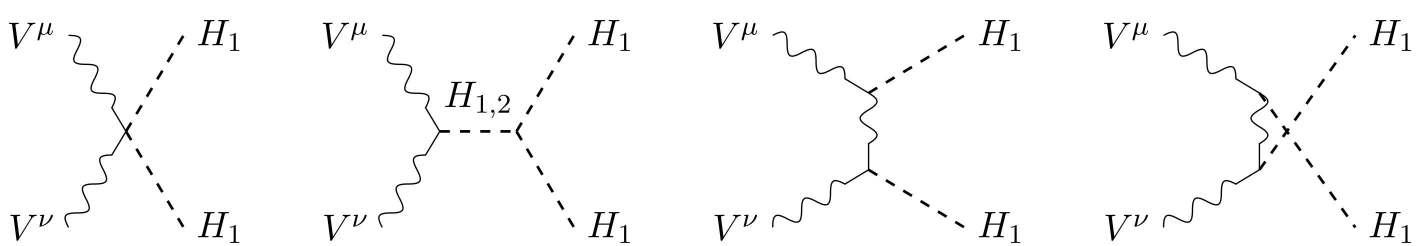
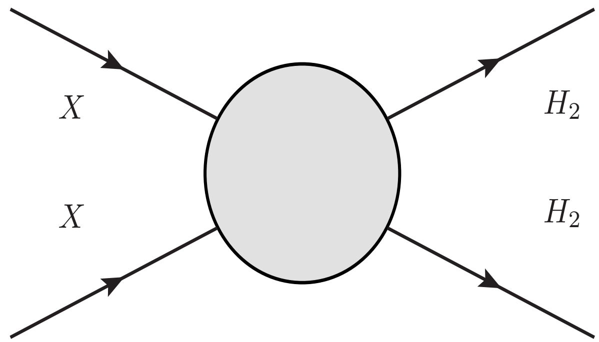


Figure 3. Dominant s/t-channel production of H_1s that decay dominantly to $b + \bar{b}$



P.Ko, Yong Tang.
arXiv:1504.03908

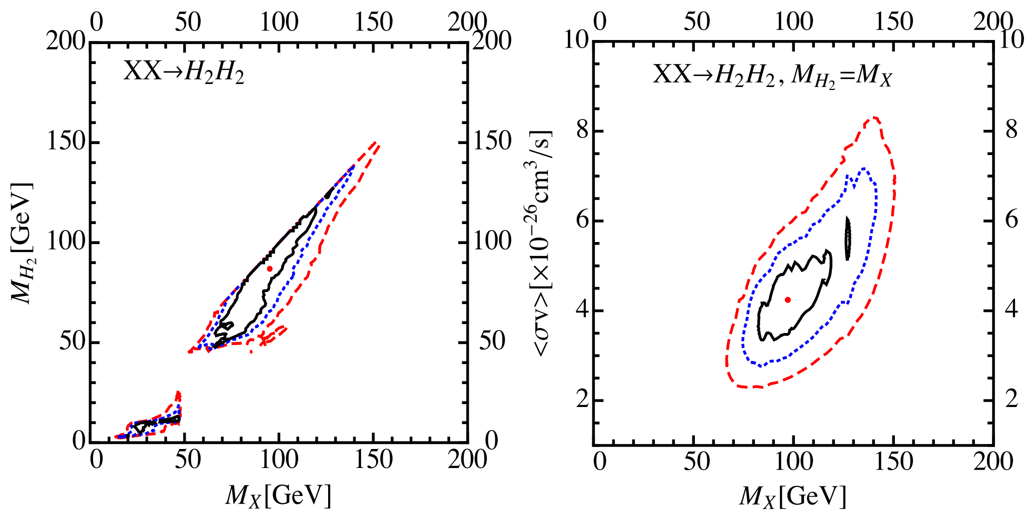


FIG. 3: The regions inside solid(black), dashed(blue) and long-dashed(red) contours correspond to 1σ, 2σ and 3σ, respectively. The red dots inside 1σ contours are the best-fit points. In the left panel, we vary freely M_X, M_{H_2} and $\langle \sigma v \rangle$. While in the right panel, we fix the mass of H_2, $M_{H_2} \simeq M_X$.

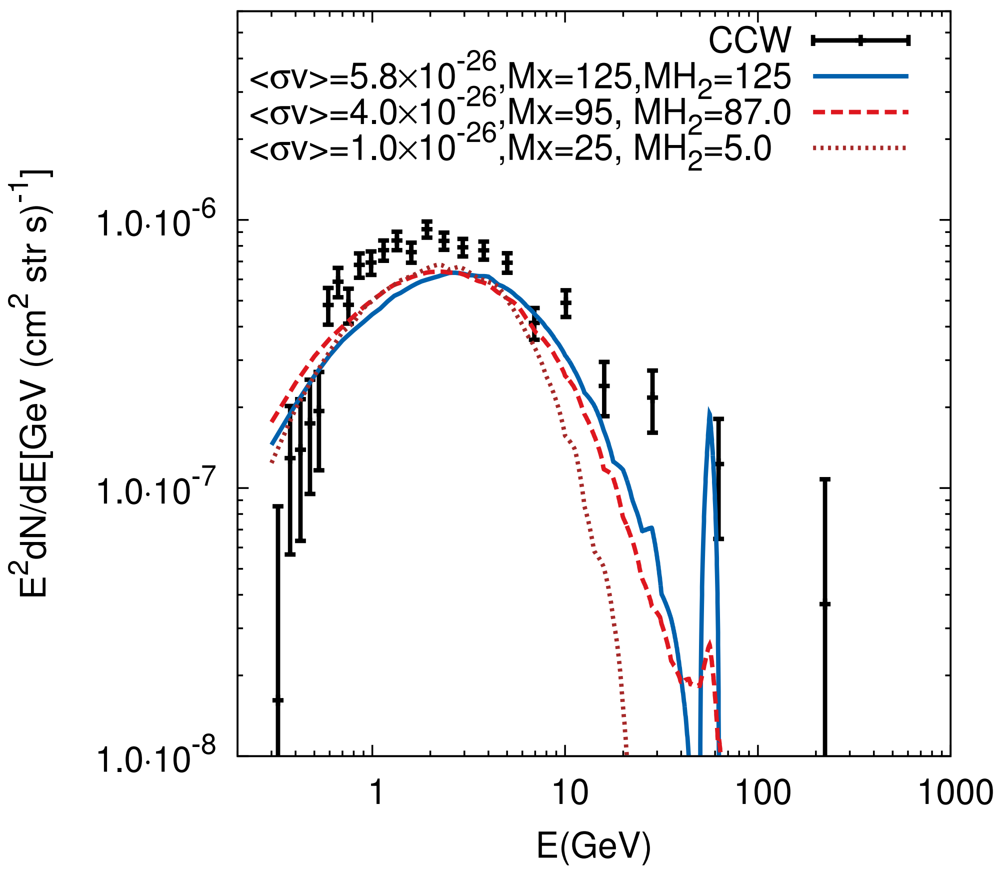


FIG. 2: Three illustrative cases for gamma-ray spectra in contrast with CCW data points [11]. All masses are in GeV unit and σv with cm^3/s. Line shape around $E \simeq M_{H_2}/2$ is due to decay modes, $H_2 \to \gamma\gamma, Z\gamma$.

Thanks to C. Weniger
for the covariant matrix

This explanation is possible only in DM models with dark gauge symmetry

P.Ko, Yong Tang. arXiv:1504.03908

Channels	Best-fit parameters	$\chi^2_{\rm min}/\text{d.o.f.}$	p-value
$XX \to H_2 H_2$ (with $M_{H_2} \neq M_X$)	$M_X \simeq 95.0\text{GeV}, M_{H_2} \simeq 86.7\text{GeV}$ $\langle \sigma v \rangle \simeq 4.0 \times 10^{-26}\text{cm}^3/\text{s}$	$22.0/21$	0.40
$XX \to H_2 H_2$ (with $M_{H_2} = M_X$)	$M_X \simeq 97.1\text{GeV}$ $\langle \sigma v \rangle \simeq 4.2 \times 10^{-26}\text{cm}^3/\text{s}$	$22.5/22$	0.43
$XX \to H_1 H_1$ (with $M_{H_1} = 125\text{GeV}$)	$M_X \simeq 125\text{GeV}$ $\langle \sigma v \rangle \simeq 5.5 \times 10^{-26}\text{cm}^3/\text{s}$	$24.8/22$	0.30
$XX \to b\bar{b}$	$M_X \simeq 49.4\text{GeV}$ $\langle \sigma v \rangle \simeq 1.75 \times 10^{-26}\text{cm}^3/\text{s}$	$24.4/22$	0.34

TABLE I: Summary table for the best fits with three different assumptions.

In Short, Dark Gauge Symmetry

- guarantees the absolute stability of weak scale DM due to unbroken (sub)group
- or guarantees its longevity due to accidental global symmetry of the underlying gauge symmetry (like baryon # in the SM)
- naturally houses DM, DR, Dark Force Carriers (dark photon, dark Higgs etc.) and interactions among them and interactions with the SM particles, resulting rich dark phenomenology
- the only issues : mass scales and coupling strengths

Models for Interacting DM-DR

- Light sterile fermion DR + Dark photon
- Nonabelian DM + DR
- (Hidden charged DM and chiral DR)

A Light Dark Photon

P.Ko, YT,1608.01083(PLB)

- Lagrangian

$$\mathcal{L} = -\frac{1}{4}V_{\mu\nu}V^{\mu\nu} + D_\mu\Phi^\dagger D^\mu\Phi + \bar{\chi}\left(i\not{D} - m_\chi\right)\chi + \bar{\psi}i\not{D}\psi$$
$$-\left(y_\chi\Phi^\dagger\bar{\chi}^c\chi + y_\psi\Phi\bar{\psi}N + h.c.\right) - V(\Phi, H),$$

- DM χ (+1), dark radiation ψ(+2), scalar(+2)
- *U(1)* symmetry (*unbroken*), massless dark photon V_μ **(Phi VEV = 0)**
- Φ is responsible for the DM relic density

$$\Omega h^2 \simeq 0.1 \times \left(\frac{y_\chi}{0.7}\right)^{-4}\left(\frac{m_\chi}{\mathrm{TeV}}\right)^2.$$

- Φ can decay into ψ and *N*.

Dark Radiation δN_{eff}

- Effective Number of Neutrinos, N_{eff}

$$\rho_R = \left[1 + N_{\rm eff} \times \frac{7}{8}\left(\frac{4}{11}\right)^{4/3}\right]\rho_\gamma,$$

$$\rho_\gamma \propto T_\gamma^4$$

- In SM cosmology, N_{eff}=3.046. Neutrinos decouple around MeV, and then freely stream.
- Cosmological bounds

Joint CMB+BBN, 95% CL preferred ranges *Planck 2015, arXiv:1502.01589*

$$N_{\rm eff} = \begin{cases} 3.11^{+0.59}_{-0.57} & \text{He+}\textit{Planck}\text{ TT+lowP,} \\ 3.14^{+0.44}_{-0.43} & \text{He+}\textit{Planck}\text{ TT+lowP+BAO,} \\ 2.99^{+0.39}_{-0.39} & \text{He+}\textit{Planck}\text{ TT,TE,EE+lowP,} \end{cases}$$

Constraint on New Physics

$$\left.\begin{aligned} N_{\rm eff} &< 3.7 \\ m_{\nu,\,\rm sterile}^{\rm eff} &< 0.52\ \text{eV} \end{aligned}\right\} \quad 95\%,\ \textit{Planck}\ \text{TT+lowP+lensing+BAO.}$$

Dark Radiation δN_{eff}

- Massless dark photon and fermion will contribute

$$\delta N_{\mathrm{eff}} = \left(\frac{8}{7}+2\right)\left[\frac{g_{*s}(T_\nu)}{g_{*s}(T^{\mathrm{dec}})}\frac{g_{*s}^D(T^{\mathrm{dec}})}{g_{*s}^D(T_D)}\right]^{\frac{4}{3}},$$

where T_ν is neutrino's temperature,

g_{*s} counts the effective number of dof for entropy density in SM,

g_{*s}^D denotes the effective number of dof being in kinetic equilibrium with V_μ.

For instance, when $T^{\mathrm{dec}} \gg m_t \simeq 173\mathrm{GeV}$ for $|\lambda_{\Phi H}| \sim 10^{-6}$, we can estimate δN_{eff} at the BBN epoch as

$$\delta N_{\mathrm{eff}} = \frac{22}{7}\left[\frac{43/4}{427/4}\frac{11}{9/2}\right]^{\frac{4}{3}} \simeq 0.53, \qquad (1)$$

$\boxed{\delta N_{\mathrm{eff}}=0.4\sim1 \text{ for relaxing tension in Hubble constant}}$

Diffusion Damping

- Dark Matter scatters with radiation, which induces new contributions in the cosmological perturbation equations,

$$\dot{\delta}_\chi = -\theta_\chi + 3\dot{\Phi},$$
$$\dot{\theta}_\chi = k^2\Psi - \mathcal{H}\theta_\chi + S^{-1}\dot{\mu}\,(\theta_\psi - \theta_\chi),$$
$$\dot{\theta}_\psi = k^2\Psi + k^2\left(\frac{1}{4}\delta_\psi - \sigma_\psi\right) - \dot{\mu}\,(\theta_\psi - \theta_\chi),$$

where dot means derivative over conformal time $d\tau \equiv dt/a$ (a is the scale factor), θ_ψ and θ_χ are velocity divergences of radiation ψ and DM χ's, k is the comoving wave number, Ψ is the gravitational potential, δ_ψ and σ_ψ are the density perturbation and the anisotropic stress potential of ψ, and $\mathcal{H} \equiv \dot{a}/a$ is the conformal Hubble parameter. Finally, the scattering rate and the density ratio are defined by $\dot{\mu} = a n_\chi \langle \sigma_{\chi\psi} c \rangle$ and $S = 3\rho_\chi/4\rho_\psi$, respectively.

Scattering Cross Section

The averaged cross section $\langle\sigma_{\chi\psi}\rangle$ can be estimated from the squared matrix element for $\chi\psi \to \chi\psi$:

$$\overline{|\mathcal{M}|^2} \equiv \frac{1}{4}\sum_{\text{pol}} |\mathcal{M}|^2 = \frac{2g_X^4}{t^2}\left[t^2 + 2st + 8m_\chi^2 E_\psi^2\right], \qquad (9)$$

where the Mandelstam variables are $t = 2E_\psi^2(\cos\theta - 1)$ and $s = m_\chi^2 + 2m_\chi E_\psi$, where θ is the scattering angle, and E_ψ is the energy of incoming ψ in the rest frame of χ. Integrated with a temperature-dependent Fermi-Dirac distribution for E_ψ, we find that $\langle\sigma_{\chi\psi}\rangle$ goes roughly as $g_X^4/(4\pi T_D^2)$.

- *In general, the cross section could have different temperature dependence, depending on the underlying particle models.*

Numerical Results

We take the central values of six parameters of ΛCDM from `Planck`,

$\Omega_b h^2 = 0.02227,$ Baryon density today

$\Omega_c h^2 = 0.1184,$ CDM density today

$100\theta_{\rm MC} = 1.04106,$ $100 \times$ approximation to r_*/D_A

$\tau = 0.067,$ Thomson scattering optical depth

$\ln\left(10^{10} A_s\right) = 3.064,$ Log power of primordial curvature perturbations

$n_s = 0.9681,$ Scalar Spectrum power-law index

which gives $\sigma_8 = 0.817$ in vanilla ΛCDM cosmology.

With the same input as above, now take

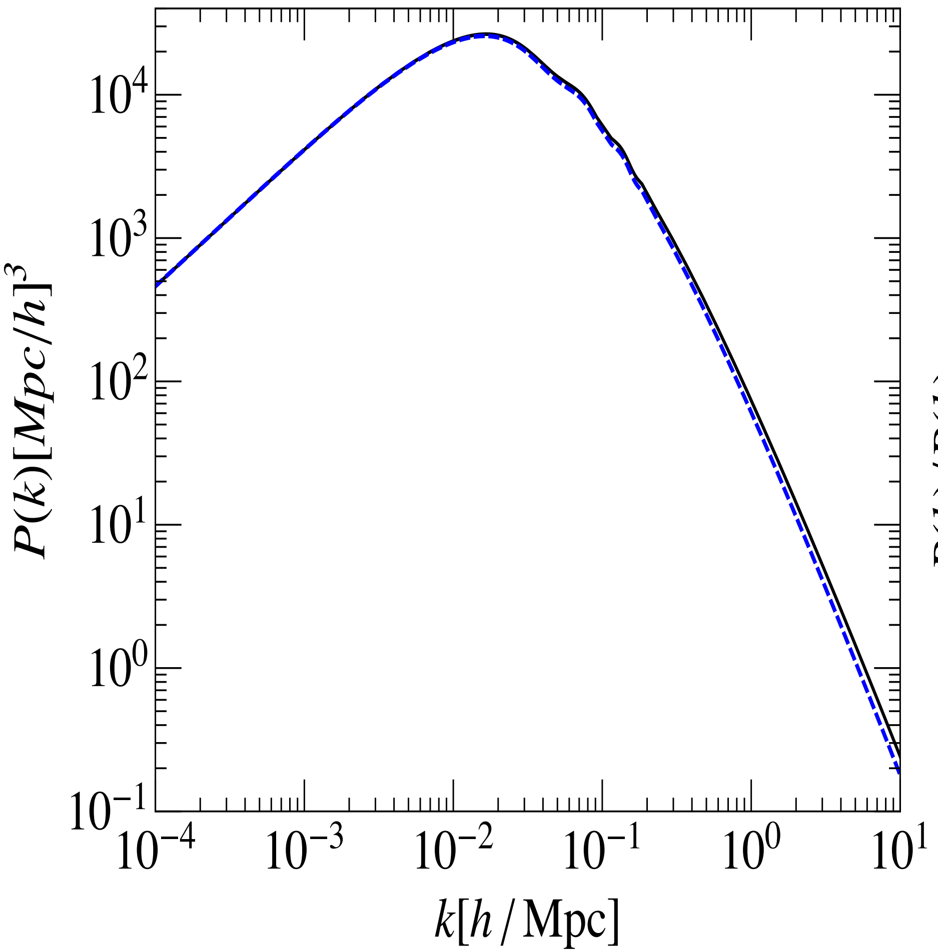
$$\delta N_{\rm eff} \simeq 0.53, m_\chi \simeq 100\text{GeV and } g_X^2 \simeq 10^{-8}$$

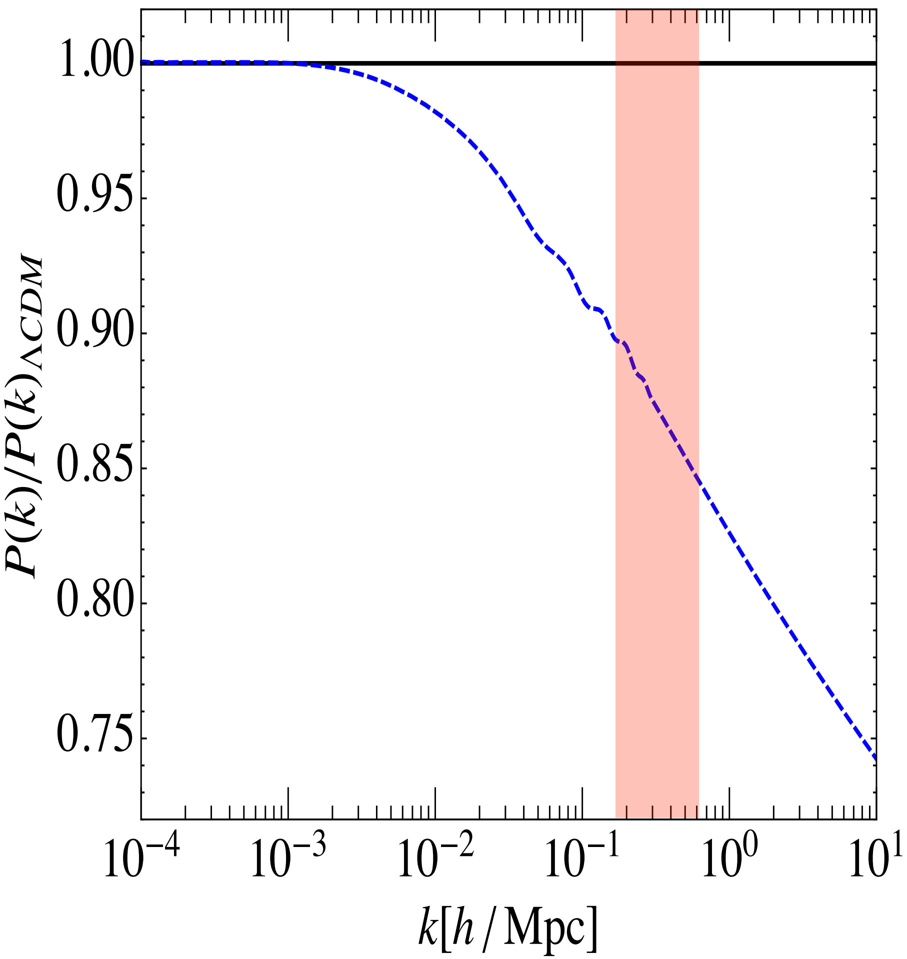
in the interacting DM case, we have $\sigma_8 \simeq 0.744$.

Modified Boltzmann code ***CLASS(Blas&Lesgourgues&Tram)***

Matter Power Spectrum

DM-DR scattering causes diffuse damping at relevant scales, resolving σ_8 problem





Results

We take the central values of six parameters of ΛCDM from `Planck` [1],

$$\Omega_b h^2 = 0.02227, \Omega_c h^2 = 0.1184, 100\theta_{\mathrm{MC}} = 1.04106,$$
$$\tau = 0.067, \ln\left(10^{10} A_s\right) = 3.064, n_s = 0.9681, \qquad (11)$$

which gives $\boxed{\sigma_8 = 0.817}$ in vanilla ΛCDM cosmology. With the same input as above, now we take $\boxed{\delta N_{\mathrm{eff}} \simeq 0.53,}$ $m_\chi \simeq 100\mathrm{GeV}$ and $g_X^2 \simeq 10^{-8}$ in the interacting DM case, we have $\boxed{\sigma_8 \simeq 0.744}$ which is much closer to the value $\sigma_8 \simeq 0.730$ given by weak lensing survey CFHTLenS [3].

Residual Non-Abelian DM&DR

P.Ko&YT, 1609.02307

- Consider *SU(N)* Yang-Mills gauge fields and a Dark Higgs field Φ

$$\mathcal{L} = -\frac{1}{4} F^a_{\mu\nu} F^{a\mu\nu} + (D_\mu \Phi)^\dagger (D^\mu \Phi) - \lambda_\phi \left(|\Phi|^2 - v_\phi^2/2 \right)^2 ,$$

- Take *SU(3)* as an example,

$$A^a_\mu t^a = \frac{1}{2} \begin{pmatrix} A^3_\mu + \frac{1}{\sqrt{3}} A^8_\mu & A^1_\mu - iA^2_\mu & A^4_\mu - iA^5_\mu \\ A^1_\mu + iA^2_\mu & -A^3_\mu + \frac{1}{\sqrt{3}} A^8_\mu & A^6_\mu - iA^7_\mu \\ A^4_\mu + iA^5_\mu & A^6_\mu + iA^7_\mu & -\frac{2}{\sqrt{3}} A^8_\mu \end{pmatrix} .$$

- *SU(3)* → *SU(2)*

$$\langle \Phi \rangle = \begin{pmatrix} 0 & 0 & \frac{v_\phi}{\sqrt{2}} \end{pmatrix}^T , \Phi = \begin{pmatrix} 0 & 0 & \frac{v_\phi + \phi(x)}{\sqrt{2}} \end{pmatrix}^T ,$$

The massive gauge bosons $A^{4,\cdots,8}$ as dark matter obtain masses,

$$m_{A^{4,5,6,7}} = \frac{1}{2} g v_\phi , \; m_{A^8} = \frac{1}{\sqrt{3}} g v_\phi ,$$

and massless gauge bosons $A^{1,2,3}_\mu$. The physical scalar ϕ can couple to $A^{4,\cdots,8}_\mu$ at tree level and to $A^{1,2,3}$ at loop level.

$SU(N) \to SU(N-1)$

- *2N-1* massive gauge bosons: Dark Matter
- *(N-1)²-1* massless gauge bosons: Dark Radiation
- mass spectrum

$$m_{A^{(N-1)^2,\dots,N^2-2}} = \frac{1}{2} g v_\phi, \; m_{A^{N^2-1}} = \frac{\sqrt{N-1}}{\sqrt{2N}} g v_\phi,$$

This can be proved by looking at the structure of f^{abc}. Divide the generators t^a into two subset,

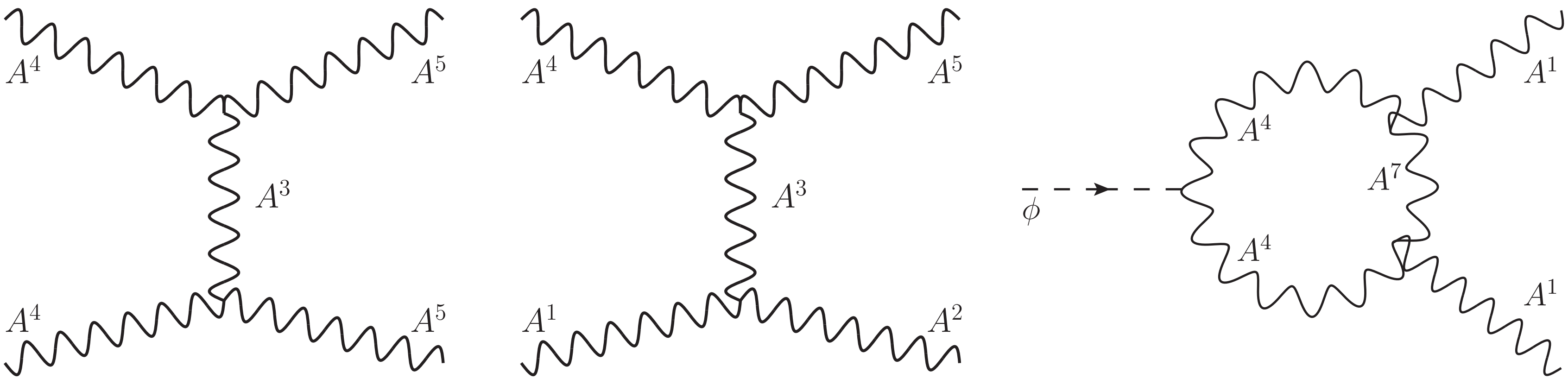
$$a \subset [1, 2, ..., (N-1)^2 - 1], a \subset [(N-1)^2, ..., N^2 - 1].$$

Since $[t^a, t^b] = i f^{abc} t^c$ for the first subset forms closed $SU(N-1)$ algebra, we have $f^{abc} = 0$ when only one of a, b and c is from the second subset. If one index is $N^2 - 1$, then other two must be among the second subset to give no vanishing f^{abc}, because t^{N^2-1} commutes with t^a from $SU(N-1)$.

Phenomenology

- Scattering and decay processes

P.Ko&**YT**, 1609.02307



- Constraints

$$\delta N_{\rm eff} = \frac{8}{7}\left[(N-1)^2 - 1\right] \times 0.055,$$

$$g^2 \lesssim \frac{T_\gamma}{T_A}\left(\frac{m_A}{M_P}\right)^{1/2} \sim 10^{-7},$$

$$\frac{m_A}{T_{\rm reh}} \sim \ln\left[\frac{\Omega_b M_P g^4}{\Omega_X m_p \eta}\right] \sim \mathcal{O}(30).$$

- ***N<6 if thermal***
- ***small coupling,***
- ***non-thermal production,***
- ***low reheating temperature***

Schmaltz et al(2015) EW charged DM

Matter Power Spectrum

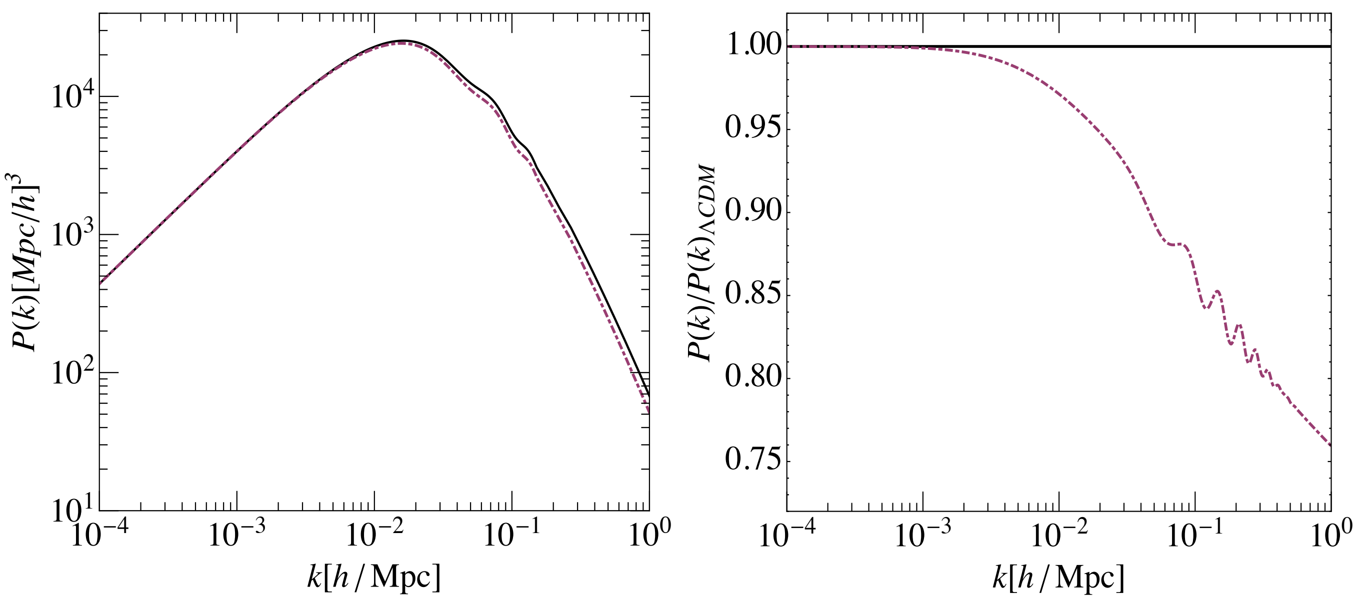


FIG. 3. Matter power spectrum $P(k)$ (left) and ratio (right) with $m_\chi \simeq 10\mathrm{TeV}$ and $g_X^2 \simeq 10^{-7}$, in comparison with ΛCDM. The black solid lines are for ΛCDM and the purple dot-dashed lines for interacting DM-DR case, with input parameters in Eq. 21. We can easily see that $P(k)$ is suppressed for modes that enter horizon at radiation-dominant era. Those little wiggles are due to the well-known baryon acoustic oscillation.

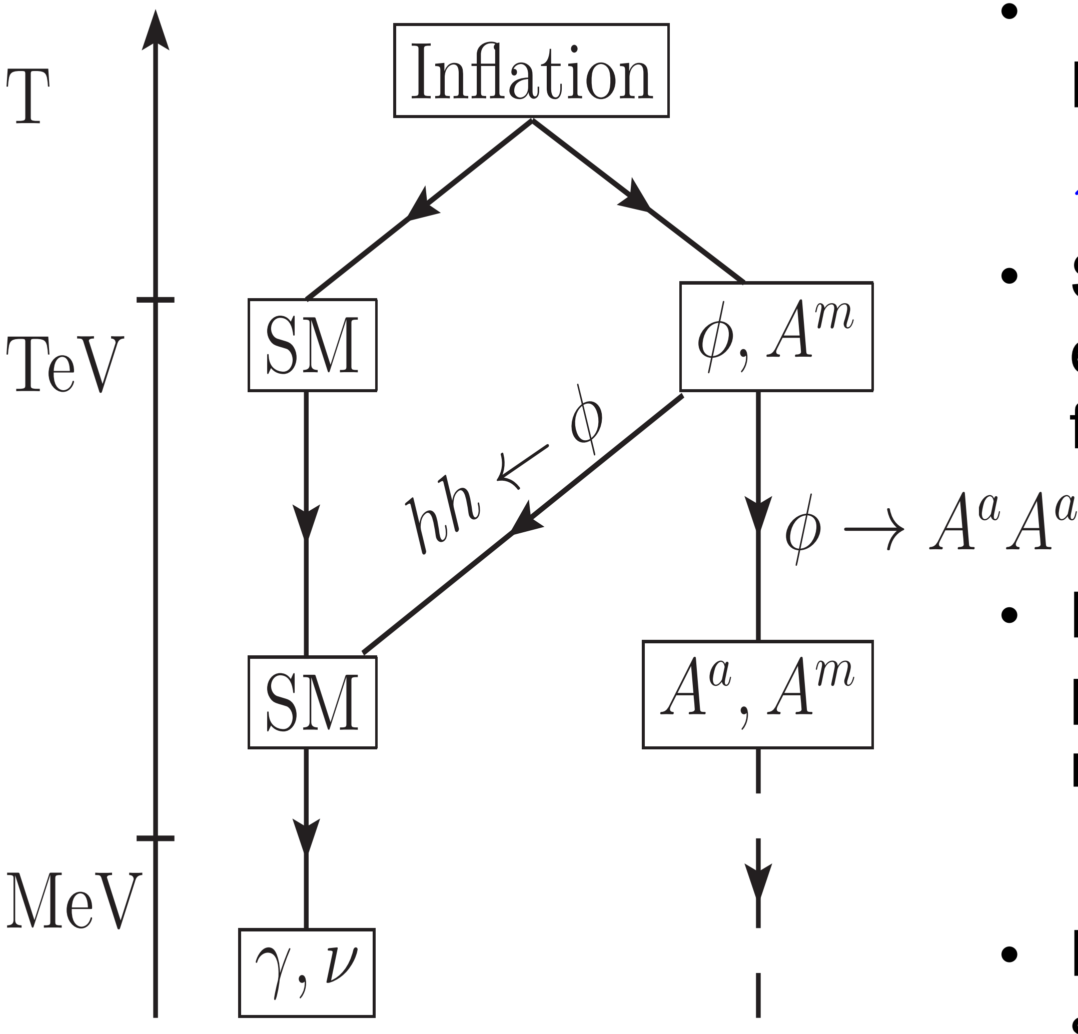
Results

$$\Omega_b h^2 = 0.02227, \Omega_c h^2 = 0.1184, 100\theta_{\rm MC} = 1.04106,$$
$$\tau = 0.067, \ln\left(10^{10} A_s\right) = 3.064, n_s = 0.9681, \tag{21}$$

and treat neutrino mass the same way as `Planck` did with $\sum m_\nu = 0.06\text{eV}$, which gives $\sigma_8 = 0.815$ in vanilla ΛCDM cosmology. Together with the same inputs as above, we take $\delta N_{\rm eff} \simeq 0.5$, $m_\chi \simeq 10\text{TeV}$ and $g_X^2 \simeq 10^{-7}$ in the interacting DM-DR case, we have $\sigma_8 \simeq 0.746$ which is much closer to the value $\sigma_8 \simeq 0.730$ given by weak lensing survey CFHTLenS [12].

- Within DM models with local dark SU(3) broken into SU(2), DM, DR and their interactions have common origin!
- And we could increase Neff, H_0 whereas making σ_8 decrease, thereby relaxing the tension between H_0 and σ_8

Thermal History



- The minimal setup with Higgs portal interaction $\lambda_{\phi H} \Phi^\dagger \Phi H^\dagger H$
- SM and DS are decoupled early, DM is produced by freeze-in mechanism
- Late time decay, entropy production due to non-relativistic decay, DR(δN_{eff})
- DM and DS scattering suppress the matter power spectrum

Summary

- We discussed some cosmological effects with *interacting* Dark Matter and Dark Radiation within DM models with dark gauge symmetries
- This scenario is motivated theoretically and also from observational tensions, H_0 and σ_8
- We present two particle physics models:
  - A massless dark photon with *unbroken U(1)* gauge symmetry
  - *Residual non-Abelian* Dark Matter and Dark Radiation
- It is possible to resolve tensions simultaneously